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Pressure balance in the multiphase ISM of cosmologically simulated disk galaxies

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ABSTRACT

Pressure balance plays a central role in models of the interstellar medium (ISM), but whether and how pressure balance is realized in a realistic multiphase ISM is not yet well understood. We address this question using a set of FIRE-2 cosmological zoom-in simulations of Milky Way-mass disk galaxies, in which a multiphase ISM is self-consistently shaped by gravity, cooling, and stellar feedback. We analyze how gravity determines the vertical pressure profile as well as how the total ISM pressure is partitioned between different phases and components (thermal, dispersion/turbulence, and bulk flows). We show that, on average and consistent with previous more idealized simulations, the total ISM pressure balances the weight of the overlying gas. Deviations from vertical pressure balance increase with increasing galactocentric radius and with decreasing averaging scale. The different phases are in rough total pressure equilibrium with one another, but with large deviations from thermal pressure equilibrium owing to kinetic support in the cold and warm phases, which dominate the total pressure near the midplane. Bulk flows (e.g., inflows and fountains) are important at a few disk scale heights, while thermal pressure from hot gas dominates at larger heights. Overall, the total midplane pressure is well-predicted by the weight of the disk gas, and we show that it also scales linearly with the star formation rate surface density (Σ_{SFR}). These results support the notion that the Kennicutt-Schmidt relation arises because Σ_{SFR} and the gas surface density (Σ_{g}) are connected via the ISM midplane pressure.

Key words: cosmology: theory – galaxies: evolution – galaxies: formation – galaxies: ISM – galaxies: star formation

1 INTRODUCTION

1.1 ISM pressure balance and its connection to the regulation of star formation in galaxies

Understanding the structure and dynamics of the interstellar medium has long been recognized as a fundamental problem in astrophysics. Since the ISM is the reservoir out of which stars form, the phase structure, dynamics, and thermodynamics of interstellar gas directly affect the process of star formation (e.g., McKee & Ostriker 2007; Krumholz 2014).

Observations also show that basic ISM properties, including its metallicity, molecular gas content, mass in gas relative to stars, and velocity dispersion vary systematically as a function of galaxy parameters such as mass and redshift (e.g., Saintonge et al. 2011; Tacconi et al. 2020). Thus, the physics of the ISM is also critical to understanding the formation and evolution of galaxies.

A principle central to essentially all models of the ISM, dating at least back to Spitzer (1956), is that of pressure balance between different phases. This principle was a core component of the two-phase ISM model proposed by Field et al. (1969) and of the three-phase model of McKee & Ostriker (1977). In these classic analytic models, the pressure of each phase is assumed to be dominated by the thermal

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component, implying that $P_{\text{therm}} = nkT$ is constant across phases. In the more modern picture, informed by a combination of observations and numerical simulations, the ISM is recognized as the host of many more complex processes. In particular, the structure and dynamics of the ISM are not only determined by thermal pressure but also by the mass, momentum, and energy stored and carried by turbulence (e.g., Elmegreen & Scalo 2004; Mac Low & Klessen 2004; Hennebelle & Falgarone 2012). Moreover, observations show (and models predict) that in many regimes star-forming galaxies drive energetic galactic winds (e.g., Heckman & Thompson 2017; Fielding et al. 2018). At their base, these galactic outflows can be viewed as flows which are not static but rather accelerated in the ISM. Thus, at the very least, a more complete understanding of the ISM should involve *dynamic* equilibrium (Kim et al. 2011; Sun et al. 2020).

Beyond the structure of the ISM, our interest in pressure balance is motivated by the role this principle may play in regulating star formation on galactic scales. Galaxies are observed to exhibit a tight correlation between their globally averaged gas surface density, Σ_g , and star formation rate surface density, Σ_{SFR} (e.g., Kennicutt 1998; Genzel et al. 2010; Kennicutt & Evans 2012). This correlation is referred to as the Kennicutt-Schmitt (KS) relation. Several analytic models have been developed in which the KS relation arises as a result of a balance between feedback and gravity (e.g., Thompson et al. 2005; Ostriker et al. 2010; Ostriker & Shetty 2011; Faucher-Giguère et al. 2013; Hayward & Hopkins 2017). Although the detailed assumptions vary between different implementations, the common basis for this class of “equilibrium” models is that the ISM pressure at the midplane must balance the weight of the overlying gas, $P_{\text{mid}} \sim \Sigma_{\text{disk}} \Sigma_g$ (where $\Sigma_{\text{disk}} = \Sigma_s + \Sigma_g$ is the total mass surface density). Assuming that $P_{\text{mid}} \propto \Sigma_{\text{SFR}}$, which can be motivated in scenarios in which the energy necessary to maintain the ISM pressure is provided either by stellar feedback or by the release of gravitational energy by inspiraling gas (e.g., Krumholz et al. 2018), this directly predicts a relation of the form $\Sigma_{\text{SFR}} \propto \Sigma_{\text{disk}} \Sigma_g$.

While the equilibrium KS models are appealing for their simplicity and ability to explain basic observations, it is not *a priori* guaranteed that the model assumptions (e.g. that a steady state is reached, that vertical pressure balance is achieved, or that pressure is correlated with the star formation rate) are realized in actual galaxies. As mentioned above, real galaxies have a dynamic, multiphase ISM shaped by a combination of gravitational, cooling, and stellar feedback processes. Additionally, in a realistic cosmological environment, galaxies continuously accrete new gas from the intergalactic medium (IGM) and frequently experience interactions (including mergers) with nearby galaxies. Moreover, many galaxies drive powerful outflows that eject gas from the ISM into the IGM (Steidel et al. 2010; Muratov et al. 2015). These complexities are typically neglected in existing analytic models (but see Hayward & Hopkins 2017, for an exception) but their effects can be investigated using simulations which include some or all of these processes. The focus of this paper is on dissecting the structure of and analyze vertical pressure balance in the ISM of simulations which are fully cosmological yet of sufficiently high resolution to capture its multiphase nature. Also presented is a brief analysis of the validity of equilibrium models for ex-

plaining the KS relation that emerge in such simulations; this will be expanded and built upon in future papers.

1.2 Previous numerical work

There has been significant previous work testing the assumption of vertical pressure balance in numerical simulations, which has been critical in establishing the present evidence in support of equilibrium models of the KS relation. However, previous analyses have relied on various assumptions or simplifications. While simplified setups have the advantage of enabling some parameters of the problem to be better controlled, it is also important to relax some of the assumptions in order to develop a more complete understanding of the physics. For example, a large number of previous studies were based on high-resolution local simulations of ~ 1 kpc patches of galactic disks with idealized boundary conditions (e.g., Shetty & Ostriker 2012; Kim et al. 2011, 2013; Kim & Ostriker 2015; Vijayan et al. 2020). In addition to neglecting interactions with the cosmological environment, it has been shown that local Cartesian box simulations cannot correctly model the development of galactic winds at large heights from the disk (Martizzi et al. 2016; Fielding et al. 2017). It is unclear whether this has implications for pressure balance at the galaxy-wind interface. Some studies of ISM pressure balance have also been carried out using global simulations of isolated galaxies (Benincasa et al. 2016; Benincasa et al. 2019b). These analyses of global galaxy models broadly supported the results of local-box simulations that the ISM tends to settle into a statistical steady state whose vertical structure is determined by a balance between pressure and gravity. However, these simulations of isolated galaxies neglected the cosmological context. Moreover, both local-box and isolate galaxy simulations previously analyzed for vertical pressure balance implemented simplified models of stellar feedback. For instance, the simulations referenced above modeled supernovae and/or photoelectric heating by FUV radiation, but did not explicitly include stellar winds, radiation pressure on dust, or model the disruption of molecular clouds by HII regions. Since the assumptions for stellar feedback can have large effects on the dynamics and phase structure of the ISM, it is important to explore different, and also more complete, sets of feedback models.

1.3 This paper

In this work, we investigate ISM pressure balance and its implications for the KS relation using the FIRE-2 simulations from the FIRE project (Hopkins et al. 2014, 2018b)¹. These are high-resolution, fully cosmological, zoom-in simulations which implement a detailed model for the ISM, star formation, and stellar feedback. The result is a dynamic, multiphase ISM self-consistently shaped by gravity, cooling, and feedback. The FIRE-2 simulations have been shown to provide a good match to several basic observed properties of galaxy populations, including the stellar mass-halo mass relation (Hopkins et al. 2014, 2018b), the mass-metallicity relation (Ma et al. 2016), galactic winds (e.g., Muratov et al. 2015; Anglés-Alcázar et al. 2017), lifetimes and properties of

¹ FIRE project website: <http://fire.northwestern.edu>

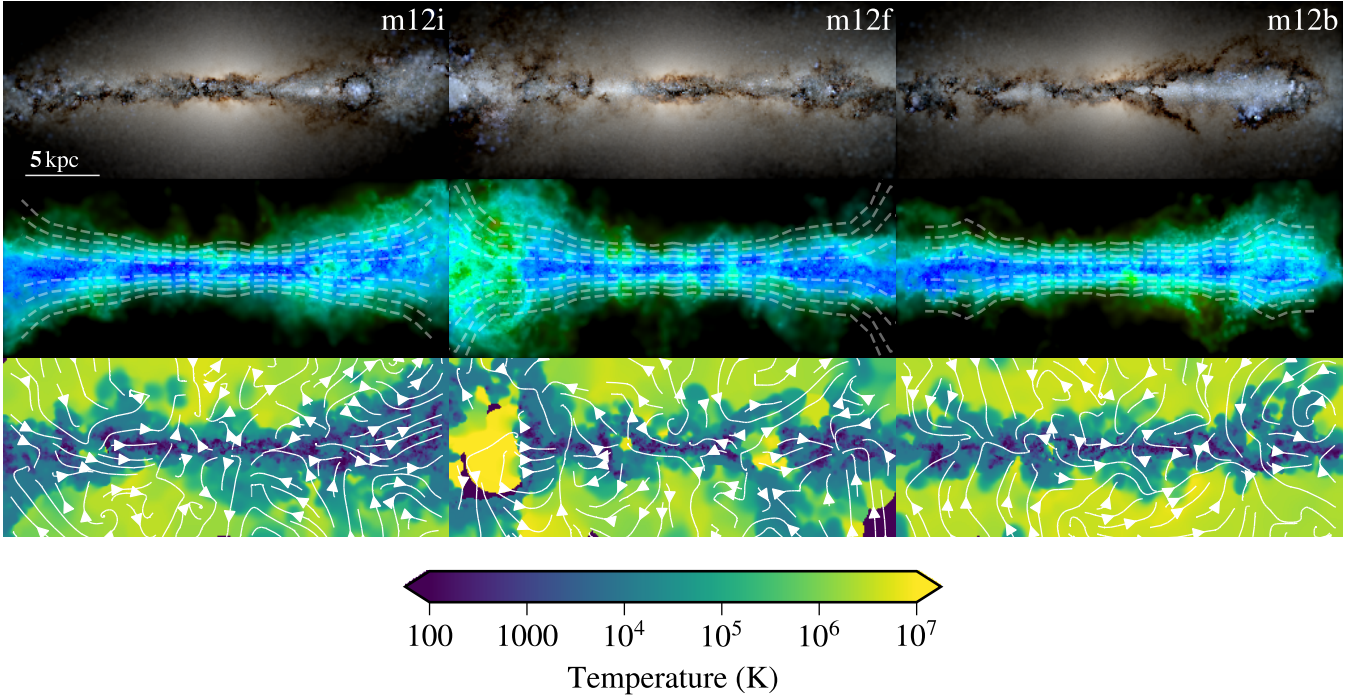


Figure 1. Renderings of Milky Way-mass galaxies from the FIRE-2 cosmological zoom-in simulations, m12i, m12f, and m12b at redshift $z = 0$ viewed edge-on. *Top:* a rendering of the mock stellar light, where pixel colors are determined by brightness in the UGR bands and including dust extinction, where emission in each band is determined by stellar age and metallicity and dust abundance is proportional to the gas metallicity. *Middle:* a projection of the gas density and temperature wherein the saturation of each pixel is proportional to the log of the column density and the hue is proportional to the log of the temperature. The result is an image where bright and blue pixels contain dense and cold gas while fuzzy and green pixels contain gas that is diffuse and hot. Dashed contours are plotted at integer values of the measured gas scale height ($\pm h$, $2h$, $3h$, and $4h$; defined in §2). *Bottom:* gas projection through a thin slice ($-250 \leq y \leq 250$ pc) colored by temperature to indicate the cold ($T \leq 10^3$ K), warm ($10^3 \leq T \leq 10^5$ K), or hot ($T \geq 10^5$ K) phase. Gas velocity streamlines are plotted in white to give an impression of the complex vertical velocity structure.

giant molecular clouds (Benincasa et al. 2019a; Guszejnov et al. 2020), and the KS relation (Orr et al. 2018).

We focus the present analysis on a set of low-redshift, Milky Way-mass simulated galaxies from the “Latte” suite of simulations (Wetzel et al. 2016). In FIRE, galaxies in this regime have large, long-lived galactic disks and relatively steady star formation rates. We systematically analyze the breakdown of the total ISM pressure into different gas phases (cold/warm/hot) and different forms of pressure (thermal/dispersion/bulk flow, which we define in §3). By comparing the pressures in different phases/forms to the weight of overlying gas we quantify the extent to which vertical pressure balance is achieved. We defer an analysis of galaxies in which star formation is highly time variable, or “bursty,” which is predicted in lower-mass galaxies and at high redshift in the FIRE simulations (e.g., Sparre et al. 2017; Faucher-Giguère 2018), to future work. We also note that this paper focuses on a theoretical analysis of the (thermo)dynamical properties of the ISM in the simulated galaxies. In particular, we do not attempt detailed comparisons with observations, which will be explored in future work. This is both because the simulation results are of interest in of themselves, and also because proper comparisons with ISM and star formation observations require detailed modeling of gas and star formation tracers. For a much more detailed assessment of how the FIRE simulations compare

with observations of the KS relation, we refer to Orr et al. (2018).

The outline of this paper is as follows. In §2 we describe the simulations analyzed in this paper and the basic properties of their galactic disks, including their Toomre Q parameter. In §3 we outline our methods for measuring pressure contributions and vertical pressure balance in our simulations. We present our main results in §4, where we quantify ISM pressure balance as a function of averaging scale, galactocentric radius, height from the midplane, gas phase, and form of pressure. We discuss our results, including in the context of equilibrium models for the origin of the KS relation, in §5. We summarize our results in §6. Appendices summarize some numerical tests and methods.

2 SIMULATIONS AND BASIC DISK PROPERTIES

2.1 Simulations

The simulations used in our analysis are FIRE-2 cosmological zoom-in simulations and are run in the Meshless Finite Mass (MFM) mode of the GIZMO² grav-

² Information about GIZMO and a public version of the code is available at: <https://www.tapir.caltech.edu/~phopkins/Site/GIZMO>

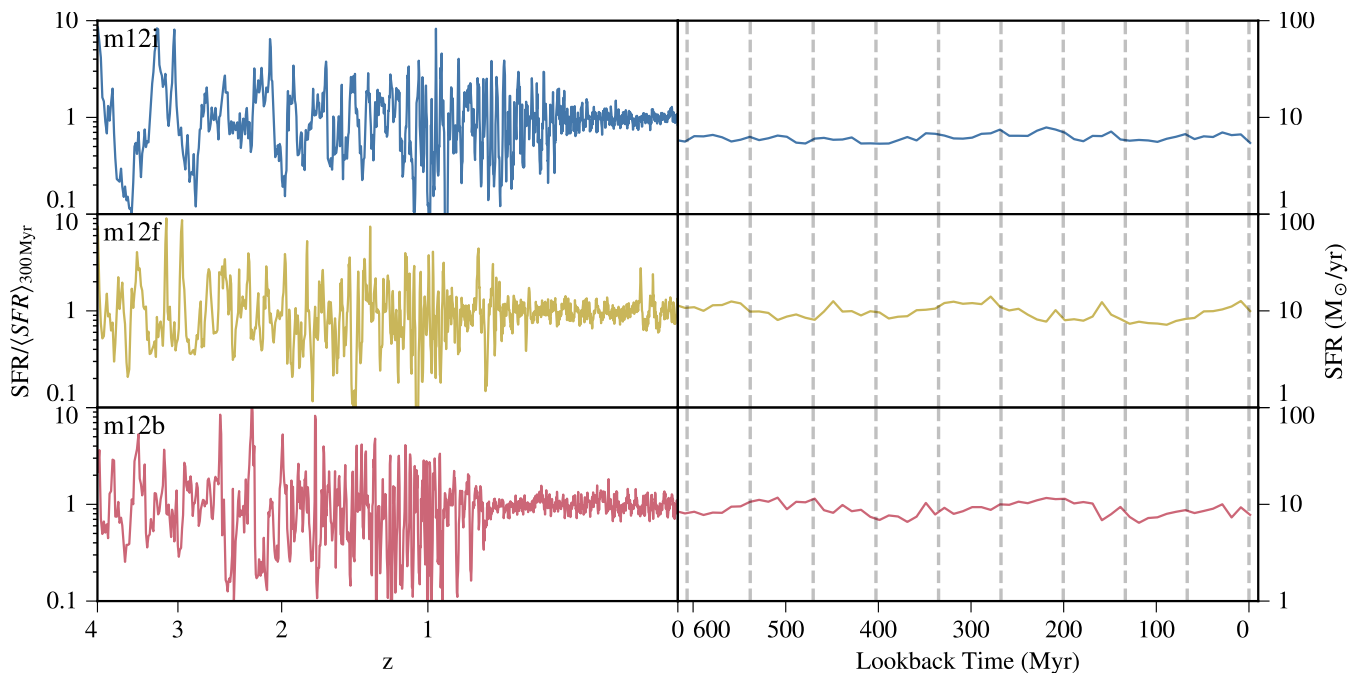


Figure 2. *Left:* Star formation histories for our three simulated galaxies, normalized by the 300 Myr running average. Although the SFR is highly variable at high redshift, these galaxies all undergo a significant transition at $z \leq 1$, following which the SFR becomes much less bursty and a stable, well-ordered, galactic disk forms. *Right:* Star formation histories for the same galaxies over the past 600 Myr before redshift $z = 0$, which is the time window on which the analysis in this paper focuses. Vertical dashed lines are plotted to show the times corresponding to the main snapshots we use for this analysis. The star formation rates are time steady (modulo minor stochastic fluctuations) over the course of the 600 Myr window we investigate.

ity+magnetohydrodynamic code. MFM is a Lagrangian, mesh-free, finite-mass method which combines advantages of traditional smooth particle hydrodynamics (SPH) and grid-based methods (for numerical details and tests, see [Hopkins 2015](#)). The FIRE-2 physics model, described in full detail in [Hopkins et al. \(2018b\)](#), includes radiative cooling for gas down to 10 K (including an approximate treatment of fine-structure metal and molecular lines). Star particles that represent simple stellar populations are formed in gas that is self-gravitating, dense ($n_{\text{H}} \geq 1000 \text{ cm}^{-3}$), and molecular (the importance of these different criteria is discussed in [Hopkins et al. 2013](#)). When all criteria for star formation are satisfied, the dense gas is converted into stars with 100% efficiency per local free-fall time. Thus, a low star formation efficiency on galactic scales is not “put in by hand” but rather emerges as a result of regulation by stellar feedback ([Hopkins et al. 2014; Orr et al. 2018](#)).³

Star particles return mass, metals, momentum, and energy into the ISM, using rates that are functions of the star particle’s age following the **STARBURST99** population synthesis model ([Leitherer et al. 1999](#)). These feedback processes include supernovae (Type II and Ia), stellar winds from O,

B, and AGB stars, photoelectric heating, photoionization, and radiation pressure. The model for radiation pressure includes both short-range and long-range components ([Hopkins et al. 2020](#)); the long-range component in particular can contribute spatially-smooth pressure supporting the overlying ISM which is not accounted for in the kinematic motions of the gas. We will return to this point when interpreting pressure balance near the midplane (§4).

In this paper we focus on three representative Milky Way-mass galaxies from the FIRE project with halo masses $\sim 10^{12} M_{\odot}$ near redshift $z = 0$: m12i, m12f, and m12b (see [Wetzel et al. 2016; Garrison-Kimmel et al. 2017, 2019](#), for detailed discussion of each of these simulations). In FIRE, these simulations form large, long-lived central disks at $z \lesssim 1$ ([Garrison-Kimmel et al. 2018](#)). The main galaxy in m12f experiences a significant merger at $z \approx 0.1$, but this occurs ≈ 600 Myr before the beginning of the time window on which we focus our analysis (see §2.2 below). The resolution of the baryons in these simulations is $m_{\text{b}} \approx 7100 M_{\odot}$ for gas cells and star particles. Thus, GMCs with masses $M_{\text{GMC}} \sim 10^6 - 10^7 M_{\odot}$ contain $\gtrsim 100 - 1000$ resolution elements. Gravitational softening lengths for gas cells are adaptive and reach values of ~ 5 pc at the average density for star formation (see Table 3 of [Hopkins et al. 2018b](#)), so massive star-forming complexes are also well resolved spatially.

Figure 1 shows $z = 0$ edge-on renderings of mock stellar emission (including dust extinction), density-temperature projections, and temperature slice-projections for each of the three galaxies. All three simulations show relatively thin

³ It is possible that processes other than stellar feedback, e.g. turbulence sourced by gravity, also play a role in the regulation of galaxy-scale star formation. However, feedback is essential since simulated galaxies based on the same physics (including gravity, cooling, turbulence, etc.) but neglecting stellar feedback experience runaway ISM collapse and SFRs that exceed observations by 1-2 orders of magnitude (e.g., [Hopkins et al. 2011](#)).

Table 1. Properties of the simulated galaxies analyzed in this paper

Name	M_{halo} ($10^{12} M_{\odot}$) ^a	$R_{*,1/2}$ (kpc) ^b	M_* ($10^{10} M_{\odot}$) ^c	f_g ^d	$\langle SFR \rangle_{600 \text{ Myr}}$ ($M_{\odot} \text{ yr}^{-1}$) ^e
m12i	1.1	2.9	6.7	0.20	6.2
m12f	1.5	4.0	8.5	0.24	9.7
m12b	1.3	2.8	8.8	0.16	9.0

^a The total dark matter and baryonic mass within R_{vir} , the virial radius defined as in [Bryan & Norman \(1998\)](#).

^b The 3D radius containing half of the total stellar mass within 15% of R_{vir} .

^c The total stellar mass within the galaxy (defined as within $5R_{*,1/2}$ cylindrically and within ± 5 kpc of the midplane).

^d $f_g = \frac{M_g}{M_* + M_g}$ with M_g the total gass mass within the galaxy.

^e Average star formation rate in the 600 Myr period before redshift $z = 0$.

($\sim 400 \text{ pc} - 1 \text{ kpc}$) disks with multiphase gas distributions. Figure 2 shows the star formation history of these three galaxies and illustrates that the star formation rates are time-steady at $z \sim 0$, despite much more bursty earlier histories.

We note that the average SFRs in the present simulations are somewhat higher than the Milky Way ($\sim 1.7 M_{\odot} \text{ yr}^{-1}$; [Bland-Hawthorn & Gerhard 2016](#)) because (1) they have higher gas masses and (2) the Milky Way itself has an SFR below the typical L_* galaxy today by a factor of 2-3 (e.g. [Licquia et al. 2015](#)). Therefore, while the simulated galaxies are hosted by Milky Way-mass dark matter halos, they are not perfect Milky Way analogues. This does not affect the validity of our results, since we are interested in the general physical processes that regulate star formation in disk galaxies and not in reproducing the detailed properties of our Galaxy.

2.2 Disk properties

We locate the main galaxy in each zoom-in simulation output using the Amiga Halo Finder (AHF, [Gill et al. 2004](#); [Knollmann & Knebe 2009](#)). We define halos using the overdensity criterion (following [Bryan & Norman 1998](#)), which corresponds to $\approx 100\times$ the critical density at $z = 0$. We define the location of the “main galaxy” to be the halo center of mass (computed using all of the dark matter, stellar, and gas components). Since we limit our analysis to times during which the disks are stable and do not experience major mergers, we find the halo center of mass to be a sufficiently accurate definition of galaxy center. In our analysis, we focus on galactocentric radii $\leq 5R_{*,1/2}$, where $R_{*,1/2}$ is the radius that contains half of the total stellar mass that lies within 15 % of the virial radius (R_{vir}). This outer radius corresponds to the data that we process our analysis pipeline; for many plots we focus on smaller radii for which the results are most relevant. Table 1 summarizes these, and other, important properties for the three simulations we use.

We define the vertical (z) axis to lie along the total angular momentum vector of the gas contained within $5R_{*,1/2}$. Since, in principle, the stellar and gas disks can be misaligned we have checked that our results are insensitive to the choice between these options. We have also confirmed that there is no significant warping of the disk in the simulation snapshots that we analyze. The $x - y$ (or $R - \phi$) plane is normal to the z -axis, and the midplane, corresponding to $z = 0$, is defined to intersect the halo center of mass. With this orientation, we measure the gas disk scale height

h as a function of cylindrical galactocentric radius R , for different annuli of width $R_{*,1/2}/3$, by fitting an exponential profile of the form $\rho(z) \propto e^{-|z|/h}$ to the vertical distribution of gas within ± 3 kpc of the midplane. We fit independent profiles both above and below the midplane and, though these heights are typically very similar, use the average of the two scale heights to define a single h that we use for the rest of the analysis.

Figure 3 catalogs important radius-dependent properties of the three galaxies in our sample including: the gas and stellar surface densities (Σ_g , Σ_s respectively), the star formation rate surface density (Σ_{SFR}), the gas and stellar scale heights (h , h_s), the circular velocity ($v_c = \sqrt{GM_{\text{tot}}(< r)}/r$), the one-dimensional velocity dispersions of the gas and stars ($\sigma_{1D,g}$, $\sigma_{1D,s}$), the mass-weighted mean gas radial and azimuthal velocities (v_R , v_{ϕ}), and Toomre Q parameters (see [Orr et al. 2018](#) for detailed comparisons of Σ_g and Σ_{SFR} in FIRE simulations to observations of the KS relation).

For all surface densities, quantities are integrated vertically between $\pm 5h$, and for all velocities, quantities are measured using resolution elements between $\pm h$. Different curves represent the profiles from different snapshots in time, separated by ≈ 60 Myr and spanning ≈ 600 Myr, and overlap strongly (except at the outermost radii), reflecting the stability of the disks during the analysis period. We repeated the analysis presented here and in the following sections with $3\times$ as many snapshots separated by ≈ 20 Myr spanning the same 600 Myr period and find that all of our time-averaged results are well converged with the 60 Myr time-spacing (see Appendix A).

The velocity dispersions σ_{1D} for both the gas and stars are defined as the average width of the velocity distributions in the r , ϕ , and z directions measured as if the distribution were Gaussian and the 84th and 16th percentiles corresponded to the $\pm\sigma$ range i.e.

$$\sigma_{1D}^2 = \frac{1}{3} \sum_{i=r,\phi,z} \left(\frac{p_{84}(v_i) - p_{16}(v_i)}{2} \right)^2. \quad (1)$$

We find that the gas 1D-velocity dispersion is roughly constant with radius at $\sim 30 \text{ km/s}$, excluding the inner-most region (which is likely affected by the presence of a stellar bulge), while the stellar 1D-velocity dispersion decreases substantially with radius. The annulus-averaged radial velocity is small, $\sim 2 - 3 \text{ km s}^{-1}$ in magnitude with both positive and negative direction, indicating there is not a strong net axisymmetric inflow of gas. The azimuthal velocity is nearly the circular velocity for all except the innermost radii

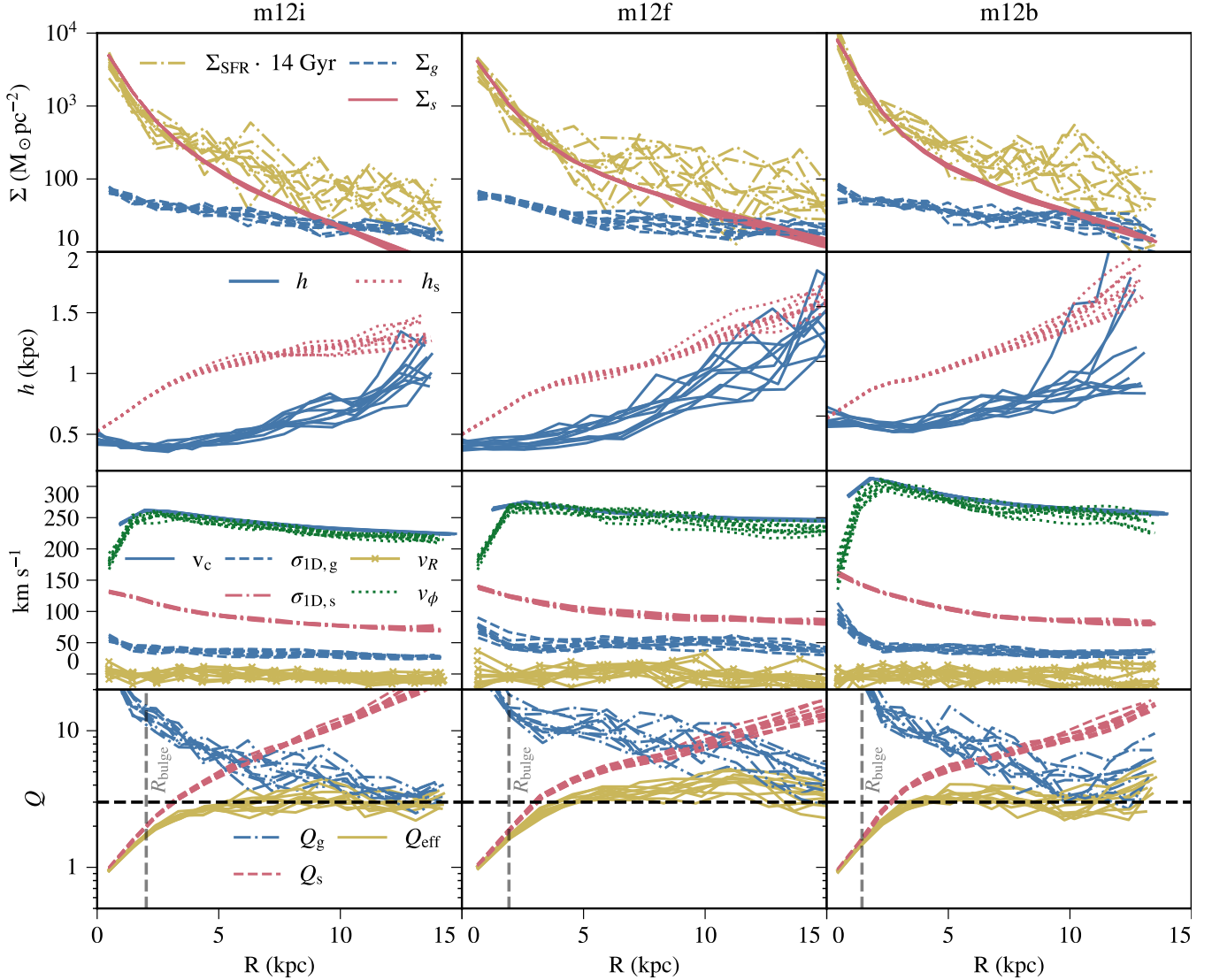


Figure 3. Disk properties vs. cylindrical radius R in annuli of width $R_{*,1/2}/3 \sim 1$ kpc in m12i, m12f, and m12b over 10 snapshots separated by ≈ 60 Myr at redshift $z \sim 0$. The bulge radius, estimated as the mass-weighted 3D radius of the counter-rotating stars, is plotted as a vertical dashed line in the bottom row. First row: surface densities of gas mass (Σ_g), stellar mass (Σ_s), and star formation rate ($\Sigma_{\text{SFR}} \cdot 14$ Gyr). Second row: gas and stellar scale heights (h and h_s respectively). Third row: circular velocity, gas and stellar velocity dispersions ($\sigma_{1D,g}$ and $\sigma_{1D,s}$ respectively), gas radial velocity (v_R), and gas azimuthal velocity (v_ϕ). Fourth row: effective Toomre Q parameter (Q_{eff}) combining Toomre parameters evaluated for gas (Q_g) and stars (Q_s) separately. Each line is a profile from a single snapshot and the relatively small variation between lines indicates the galaxies are nearly in steady state.

(see Wellons et al. 2020, for a discussion of deviations of v_ϕ from v_c , and their implications, in the FIRE simulations).

where Q_g and Q_s are defined for the gas and stars separately:

$$Q_g = \frac{\Omega \sigma_{1D,g}}{\pi G \Sigma_g} \quad \text{and} \quad Q_s = \frac{\Omega \sigma_{1D,s}}{\pi G \Sigma_s}, \quad (3)$$

2.3 Toomre Q

We also measure the the Toomre Q parameter in annuli, a proxy for the stability of a rotating disk against gravitational collapse (e.g. Toomre 1964; Elmegreen 2011). To characterize our disks, which include both gas and stars, we follow Wang & Silk (1994) who define an “effective” Q that includes the effects of a two-component disk (gas+stars):

$$Q_{\text{eff}}^{-1} = Q_s^{-1} + Q_g^{-1}, \quad (2)$$

and Ω is the orbital frequency of the system as a whole.

The bottom panel in Figure 3 shows the effective Toomre Q parameters as a function of radius in our three simulations. We find that Q_{eff} is ~ 3 for radii that are outside the influence of the galactic bulge and increases to ~ 10 for radii greater than 15 kpc (not shown), where there is little star formation, for all three simulations. We return to, and interpret, this result in §5.

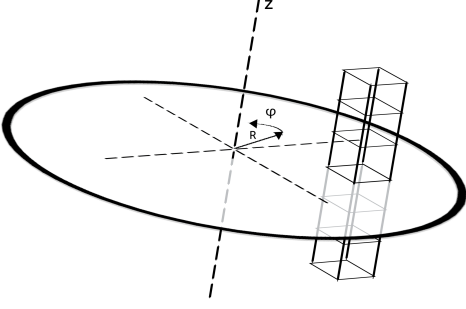


Figure 4. A schematic of the analysis setup to illustrate the orientation of our coordinate system. The z -direction is defined to lie along the axis of the gas angular momentum vector, the cylindrical R direction increases outward from the barycenter of the dark matter halo, and the ϕ direction increases in the direction of galactic rotation. An example column of “slabs” in which we measure quantities are drawn at an arbitrary position in the disk.

3 EQUATIONS OF VERTICAL PRESSURE BALANCE

3.1 Definition of ‘columns’ and ‘slabs’

Here we define a framework within which we measure the physical properties of different ISM phases, with the goal of testing whether (and how) vertical pressure balance is realized in different regions of the galaxy. To do this, we subdivide the galaxy within $5R_{*,1/2}$ into square $x-y$ “patches” (or apertures) of size $\ell_{\text{patch}} \times \ell_{\text{patch}}$, where $\ell_{\text{patch}} = 100$ pc, 300 pc, 1 kpc, and 3 kpc is varied to study the dependence of the results on averaging scale. These values for ℓ_{patch} are representative of some spatially-resolved galaxy observations (e.g. Leroy et al. 2008; Gallagher et al. 2018; Sun et al. 2020); we do not expect any of our results to depend significantly on the exact shape of the patches used to define the analysis. Each patch is assigned a scale height based on the galactocentric radius of the patch center (as described in §2.2 above). Each patch is then subdivided vertically into “columns” oriented parallel to the z -axis of 400 “slabs” of equal thickness from $-20h \rightarrow 20h$. Thus, in each column there are 200 slabs above the midplane and 200 slabs below the midplane, each of width $h/10$. The overall geometry of patches, columns, and slabs thus defined is illustrated in Figure 4.

3.2 Generalized ‘dynamic’ pressure balance

In cosmological simulations, the ISM is not necessarily in *hydrostatic* balance at all times and locations because of (i) density and velocity fluctuations; and (ii) bulk flows such as fountain flows above (or below) the disk plane that can contribute significantly to supporting the overlying gas. There can also be external perturbations, due to e.g. satellite galaxies. We therefore derive more general relationships between averaged kinematic/dynamic/thermodynamic galactic properties implied by the general momentum (Euler) fluid equation:

$$\left(\frac{\partial}{\partial t} + \vec{u} \cdot \vec{\nabla} \right) \vec{u} = -\frac{1}{\rho} \vec{\nabla} P + \vec{g} + \vec{S}, \quad (4)$$

where \vec{u} is the velocity field, P is the total effective pressure, and \vec{g} is the gravitational acceleration (due to the total mass in the system). The \vec{S} term on the right-hand side corresponds to momentum injection (expressed in units of acceleration), such as by stellar feedback. In many circumstances, we may expect that the average momentum injection rate is locally zero, e.g. because the vector sum of momentum injected by a point-like supernova explosion is zero by construction in our subgrid model (Hopkins et al. 2018a). However, some momentum injection is explicitly non-local (including the long-range radiation pressure in the FIRE model) and even nominally “local” feedback injection can locally average to a non-zero net momentum because, numerically, the momentum is injected over a finite volume consisting of ~ 32 neighboring MFM cells. We thus keep track of the momentum source term \vec{S} in the derivation below.

Next, we note that our primary interest is pressure support on scales $z \sim h$, where $|z| \ll r$ for thin disks. On these scales, we assume that the system in each slab is well approximated by a steady state with plane-parallel geometry, i.e. we only consider the z -component of Equation (4). With these approximations, and after multiplying both sides by ρ , the z -equation can be written in cylindrical coordinates as:

$$\rho \left(\frac{\partial u_z}{\partial t} + u_R \frac{\partial u_z}{\partial R} + \frac{u_\phi}{R} \frac{\partial u_z}{\partial \phi} + u_z \frac{\partial u_z}{\partial z} \right) = -\frac{\partial P}{\partial z} + \rho g_z + \rho S_z. \quad (5)$$

Let us now focus on the $u_z \partial u_z / \partial z$ term on the left. Using the mass continuity equation, $\partial \rho / \partial t + \vec{\nabla} \cdot (\rho \vec{u}) = 0$, it is straightforward to verify that

$$\rho u_z \frac{\partial u_z}{\partial z} = u_z \frac{\partial \rho}{\partial t} + u_z \frac{\partial (\rho u_R)}{\partial R} + \frac{\rho u_R u_z}{R} + \frac{u_z}{R} \frac{\partial (\rho u_\phi)}{\partial \phi} + \frac{\partial (\rho u_z^2)}{\partial z} \quad (6)$$

so that, after combining terms with like partial derivatives, Equation (5) can be expressed as:

$$\frac{\partial (\rho u_z)}{\partial t} + \frac{\partial (\rho u_R u_z)}{\partial R} + \frac{\rho u_R u_z}{R} + \frac{1}{R} \frac{\partial (\rho u_\phi u_z)}{\partial \phi} + \frac{\partial (\rho u_z^2 + P)}{\partial z} = \rho g_z + \rho S_z \quad (7)$$

Since we focus our analysis on time periods during which the simulated disks are long-lived, with approximately constant properties, we assume that the disks are well-described by a statistical steady state. Under this assumption, the $\partial/\partial t = 0$ on the left-hand side of Equation (7) can be neglected after time averaging:⁴

⁴ Note that upon time averaging between an initial time t_i and a final time t_f , $\langle \partial(\rho u_z) / \partial t \rangle \propto \int_{t_i}^{t_f} dt' \partial(\rho u_z) / \partial t' = \Delta(\rho u_z)_{t_i}^{t_f}$. Even in a statistical steady state, the instantaneous value of ρu_z can vary significantly from t_i to t_f within a single column. However, after averaging over many independent columns, the fluctuations should cancel. We show in Appendix A that our results are indeed insensitive to this term.

$$\begin{aligned}
& \left\langle \frac{\partial(\rho u_R u_z)}{\partial R} \right\rangle + \left\langle \frac{\rho u_R u_z}{R} \right\rangle \\
& \left\langle \frac{1}{R} \frac{\partial(\rho u_\phi u_z)}{\partial \phi} \right\rangle + \left\langle \frac{\partial(\rho u_z^2 + P)}{\partial z} \right\rangle \\
& = \langle \rho g_z \rangle + \langle \rho S_z \rangle
\end{aligned} \tag{8}$$

Later in the paper, we plot results up to $|z| = 10h$ to show how the flow develops as the disk joins the CGM. At these large $|z|$, the plane-parallel geometry assumption can break down significantly and so we do not necessarily expect Equation (8) to be accurately satisfied.

We then *define* the rest frame of each column such that it comoves with an annulus at radius R . Specifically, we give each column an azimuthal velocity equal to the (mass-weighted) average disk gas azimuthal velocity at R and a radial velocity equal to the average radial velocity at the same radius in annuli of width $R_{*,1/2}/3 \sim 1$ kpc (the same annuli used in §2 and Figure 3). Because of this choice, the column rest frames are generally not inertial and the Euler equation should include terms corresponding to fictitious centrifugal and Coriolis forces. However, in plane-parallel geometry these fictitious terms act purely in the $x - y$ plane and so they do not modify Equation (5). In the rest frame of any given column, we therefore have $\langle u_R \rangle = 0$ and $\langle u_\phi \rangle = 0$ by construction.

Our final assumption is that the R and ϕ components of the velocity field are statistically independent of the z component and the density field ρ , so that the time averages of terms involving products of R and z velocity components, or ϕ and z velocity components, factor out⁵:

$$\begin{aligned}
\left\langle \frac{\partial(\rho u_R u_z)}{\partial R} \right\rangle &= \frac{\partial}{\partial R} \langle u_R \rangle \langle \rho u_z \rangle = 0, \\
\left\langle \frac{\rho u_R u_z}{R} \right\rangle &= \frac{1}{R} \langle u_R \rangle \langle \rho u_z \rangle = 0, \\
\left\langle \frac{1}{R} \frac{\partial(\rho u_\phi u_z)}{\partial \phi} \right\rangle &= \frac{1}{R} \frac{\partial}{\partial \phi} \langle u_\phi \rangle \langle \rho u_z \rangle = 0.
\end{aligned} \tag{9}$$

This leaves us with the following simplified equation:

$$\left\langle \frac{\partial(P + \rho u_z^2)}{\partial z} \right\rangle = \langle \rho g_z \rangle + \langle \rho S_z \rangle \tag{10}$$

We now take a coarse-grained view by averaging both sides of Equation (10) over the volume of a slab. The superscript ‘ s ’ distinguishes slab-averaged quantities from their continuously-defined values. There is an important subtlety in the definitions that follow. Although the volume average of Equation (10) is evaluated without any weighting, we define several of the slab-averaged quantities as weighted by mass, e.g. $u_z^s = \langle u_z \rangle_M$, $g_z^s = \langle g_z \rangle_M$, $S_z^s = \langle S_z \rangle_M$, ..., where $\langle x \rangle_M = (\int dV \rho x) / m_{\text{slab}}$. The slab density is defined as $\rho^s = m_{\text{slab}} / V_{\text{slab}}$. Since the different terms in Equation (10) involve a factor of density ρ , this yields simple (but fully rigorous) expected relationships between slab-averaged quantities. In this coarse-grained view, P^s is the sum of the

⁵ This is not strictly guaranteed but is a reasonable simplifying assumption given the complex (in some respects turbulent) dynamics of the ISM.

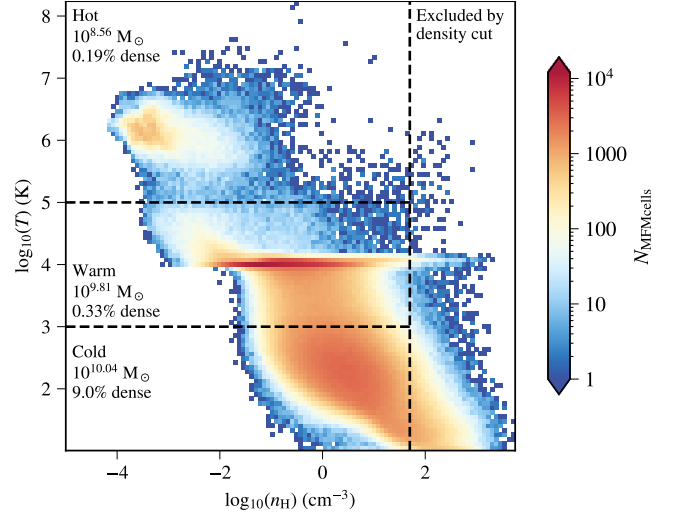


Figure 5. A temperature-density diagram of the particles within $5R_{*,1/2}$ cylindrically and within ± 20 kpc from the midplane vertically in *m12i* at redshift $z = 0$. Horizontal lines are plotted at $T = 10^3$ K and $T = 10^5$ K to denote the warm-cold and warm-hot boundaries respectively. Likewise, a vertical line is plotted at a density of 50 cm^{-3} to denote the cut above which we excise dense gas so as to exclude clumps that do not participate in vertical pressure balance. The percentage of MFM cells in each phase that are classified as “dense” (and thus excluded), along with the total mass in each phase, is annotated on the plot.

usual thermal term and a “dispersion” term (which is often referred to as the “turbulent pressure,” but which can contain motions which are not strictly turbulent) corresponding to the kinetic energy of flows that are resolved by the simulation but on scales $< \ell_{\text{patch}}$: $P^s = P_{\text{therm}}^s + P_{\text{disp}}^s$, where $P_{\text{therm}}^s = (3/5)\rho^s c_s^2$, $P_{\text{disp}}^s = \rho^s \sigma_z^2$, c_s^2 is the mass-averaged sound speed squared,⁶ and $\sigma_z^2 = \langle (u_z - u_z^s)^2 \rangle_M$. Finally, we define a “bulk flow” pressure $P_{\text{bulk}}^s \equiv \rho^s u_z^s{}^2$, which accounts for momentum transfer owing to the net flow through a slab.

After these definitions, we have:⁷

$$\left\langle \frac{\partial P_{\text{tot}}^s}{\partial z} \right\rangle = \langle \rho^s g_z^s \rangle + \langle \rho^s S_z^s \rangle, \tag{11}$$

where

$$P_{\text{tot}}^s = P_{\text{therm}}^s + P_{\text{disp}}^s + P_{\text{bulk}}^s. \tag{12}$$

Note that, since g_z^s and S_z^s are defined as mass-weighted averages over slabs, it is the case that $[\rho g_z]^s = \rho^s g_z^s$ and $[\rho S_z]^s = \rho^s S_z^s$, a fact used on the right-hand side of Equation (11).

We can then integrate each side from height z above

⁶ Note that $P_{\text{therm}}^s = \rho^s \langle (v_{\text{th},z} - u_z^s)^2 \rangle_M = (3/5)\rho^s c_s^2$ for a monatomic gas with adiabatic index $\gamma = 5/3$, where $v_{\text{th},z}$ is the z component of thermal velocities.

⁷ Since eq. (10) involves a partial derivative with respect to z on the left-hand side, which interacts with the volume average, the slab-averaged version should involve z “boundary terms,” similar to the $\Delta \rho u_z$ that we neglected when taking the time average of eq. (7). In this case, the boundary terms cancel exactly in the limiting of infinitely thin slabs, and we show in Appendix A that our results are also insensitive to these boundary terms.

the midplane to ∞ . We assume that $P_{\text{tot}}(\infty) = 0$, which is a good assumption so long as the midplane pressure is much larger than the pressure out in the CGM. Integrating the left-hand side gives the slab-averaged pressure $P_{\text{tot}}^s(z)$, while the right-hand side gives the weight per unit area of the gas above the slab at height z , if $z > 0$:

$$W_{\text{tot}}^s(z) \equiv - \int_z^\infty dz' \rho^s(z') g_z^s(z'). \quad (13)$$

with an analogous integral with bounds $-\infty \rightarrow z$ if $z < 0$ (below the midplane). Neglecting the momentum source term, as is commonly done in similar analyses (e.g., Kim & Ostriker 2015), the expression for steady state balance in Equation (11) is then equivalent to $\langle P_{\text{tot}}^s(z) \rangle = \langle W_{\text{tot}}^s(z) \rangle$ for all z . In addition to testing this result against the simulations, in this work we analyze how P_{tot} is partitioned into thermal, dispersion, and bulk flow, as a function of R and z .

In practice, to compute slab-averaged quantities, we smoothly deposit gas properties onto the grid of slabs in an SPH-like way (see Appendix B). To evaluate W_{tot} , we use `pykdgrav`⁸ to compute the gravitational acceleration in the z -direction, taking into account all the particles within the virial radius of the main halo. To minimize noise, we average four evaluations of $g_z(z)$, each at a point $\ell_{\text{patch}}/4$ away from the patch center and along a different cardinal direction. This procedure effectively replaces g_z^s (defined as a mass-weighted average) by a slab-volume averaged version; we expect this to be a minor effect as the gravitational acceleration is dominated by long-range forces except in the densest gas, which is excluded from the analysis as described in the next section.

3.3 Partial vs. local pressures in different phases

To understand how pressure support is partitioned in a multiphase ISM, we also divide the gas in the galaxy into three phases according to its temperature: cold ($T < 10^3$ K), warm ($10^3 < T < 10^5$ K), and hot ($T > 10^5$ K). In addition to this temperature cut, we exclude all gas denser than 50 cm^{-3} to excise very dense clumps that are largely decoupled from the volume-filling flow and effectively do not participate in vertical pressure balance. We find that our results are insensitive to density cuts above 10 cm^{-3} , below which substantial amounts of diffuse cold and warm gas are erroneously excised. Figure 5 shows a temperature-density diagram illustrating the different gas phases in one of our simulations.

In simplified models, such as the classic three-phase model of the ISM (e.g. McKee & Ostriker 1977), the different phases are assumed to have the same “local” pressure, i.e.

$$\tilde{P}_{\text{tot}}^{\text{cold}} \approx \tilde{P}_{\text{tot}}^{\text{warm}} \approx \tilde{P}_{\text{tot}}^{\text{hot}}, \quad (14)$$

where here we use the term “local” to refer to the pressure (\propto energy density) in the volume *occupied by a given phase*. We contrast this with the “partial” pressure, which we define as the volume weighted contribution to the total volume

integrated pressure. The partial and local pressures for each phase are then related by the volume-filling fraction of a phase, $f_{\text{phase}} = V_{\text{phase}}/V_{\text{slab}}$.

$$P^{\text{phase}} = f_{\text{phase}} \tilde{P}^{\text{phase}}. \quad (15)$$

The total ISM pressure can then be written as the sum of the partial pressures from each phase.

$$P_{\text{tot}} = P_{\text{tot}}^{\text{cold}} + P_{\text{tot}}^{\text{warm}} + P_{\text{tot}}^{\text{hot}} \quad (16)$$

In the rest of this paper, we use the “tilde” notation as above to distinguish between local vs. partial pressures.

4 VERTICAL PRESSURE BALANCE RESULTS

In this section we measure whether (and how) gas achieves vertical pressure balance in different regions of the galaxy.

4.1 Pressure-to-weight ratio maps and scale dependence

Figure 6 maps the total pressure-to-weight ratio, $P_{\text{tot}}/W_{\text{tot}}$, in a single $z = 0$ snapshot of m12i over the face of the disk as a function of ℓ_{patch} (in different panels). Each patch is colored according to the median $_{|z| \leq h} \{P_{\text{tot}}(z)/W_{\text{tot}}(z)\}$, i.e. the median pressure-to-weight ratio of slabs within $\pm h$ of the midplane. Face-on renderings using the same visualization methodology as the top panels of Figure 1 are included on the left to illustrate, e.g., the locations of spiral structure in the disk. Contours of $\Sigma_g = 10 \text{ M}_\odot/\text{pc}^2$ are reproduced in all panels to facilitate comparison between the results for different ℓ_{patch} values. The figure shows that, even for averaging scales $\ell_{\text{patch}} = 1 \text{ kpc}$, $P_{\text{tot}}/W_{\text{tot}}$ fluctuates substantially from place to place within a given galaxy. Although there are some apparent correlations between gas surface density contours and the pressure ratio (e.g., the inner galaxy, where Σ_g is large, generally has $P_{\text{tot}}/W_{\text{tot}} < 1$), *overall the pattern is complex and spiral features can correspond to either “over-pressurized” or “under-pressurized” regions*. Outside the central region, gas near the midplane typically has a pressure-to-weight ratio ≤ 1 .

The fluctuations shown in Figure 6 highlight the extent to which local box simulations, which typically cover disk patches of size $\leq 1 \text{ kpc}$, cannot capture the full range of conditions occurring throughout a galaxy’s ISM. Next, we analyze more quantitatively how the pressure-to-weight ratio varies as a function of galactocentric radius and height from the midplane.

4.2 Pressure-to-weight ratio vs. R and z

The top row of Figure 7 shows the height-median pressure ratio within $\pm h$, more precisely defined as median $_{|z| \leq h} \{\text{mean}_{R,\phi,t}(P_{\text{tot}})/\text{mean}_{R,\phi,t}(W_{\text{tot}})\}$, for different values of ℓ_{patch} and the three simulations analyzed in this paper. This formula indicates that, for fixed z/h , we first average over patches within a given annulus (R, ϕ), including patches from snapshots at different times (t) in each annulus. Then, for each annulus, we report the median for all slabs with $|z| \leq h$. Overall, the height-median pressure ratio within $\pm h$ is between 0.5-1 for all three simulations

⁸ `pykdgrav` is a Python package for efficiently computing N -body gravitational accelerations (and potentials) using a k-d tree. See <https://github.com/mikegrudic/pykdgrav/>.

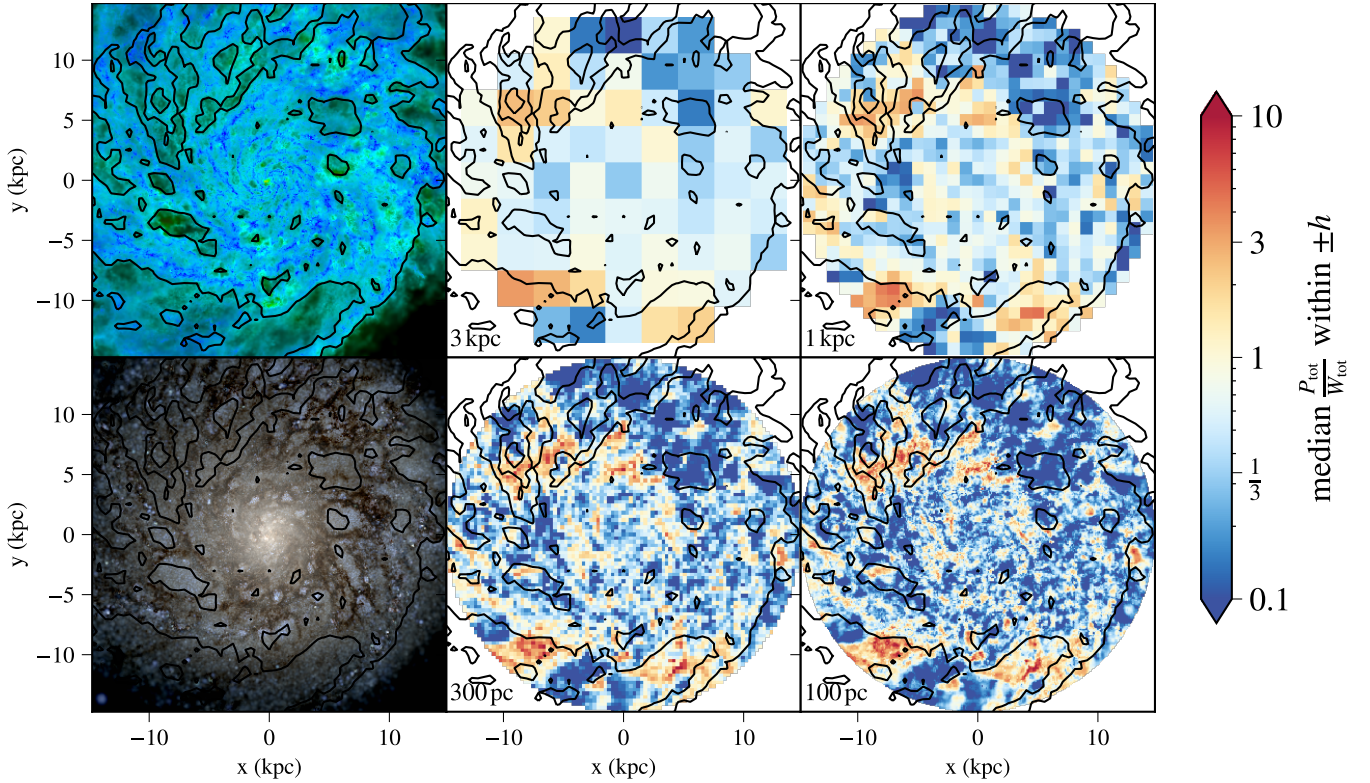


Figure 6. The two columns on the right show the median pressure-to-weight ratio within $|z| \leq h$, including all phases, at redshift $z = 0$ in m12i. Different panels show different patch sizes ℓ_{patch} , denoted in the bottom left corner of each panel. Face-on renderings of the gas (top left) and stars (bottom left) analogous to edge-on renderings in Figure 1 are provided as a reference for spiral and other high-density features. The contours correspond to gas surface density $\geq 10 \text{ M}_{\odot}/\text{pc}^2$ and are intended to aid comparison between panels. In general, deviations from vertical pressure balance increase with decreasing patch size (from top middle to bottom right) owing to increasing scatter. The majority of the volume is occupied by under-dense and under-pressurized gas whereas over-pressurized gas correlates with heating by stellar feedback around actively star-forming regions.

at all radii plotted and systematically increases with radius, independent of the averaging scale ℓ_{patch} .

The bottom row of Figure 7 shows a measure of scatter in the pressure ratio vs. radius and ℓ_{patch} , which we define as $\text{median}_{|z| \leq h} \{ \text{IQR}_{r, \phi, t}(P_{\text{tot}}/W_{\text{tot}}) \}$, where IQR is the interquartile range (the distance between the 25th and 75th percentiles). This may be viewed as the “height-median of the patch-to-patch scatter within $\pm h$.” We find that the scatter increases both as a function of radius and as ℓ_{patch} is decreased for all simulations. These trends are, at least qualitatively, consistent with scale-dependent sampling effects (e.g., Torrey et al. 2017; Hani et al. 2020). In particular, fewer star-forming complexes are contained within each patch on average with increasing radius and decreasing patch size.

In addition to measuring the height-median within $\pm h$, we also examine vertical profiles of the pressure ratio for three different, large radial annuli ($R = 0 - 3$, $3 - 6$, $6 - 9$ kpc) in Figure 8. In this figure, the vertical profiles show the ratio of means, i.e. $\langle P_{\text{tot}} \rangle / \langle W_{\text{tot}} \rangle(z) = [\text{mean}_{r, \phi, t}(P_{\text{tot}}) / \text{mean}_{r, \phi, t}(W_{\text{tot}})](z)$, which the derivation in §3.2 shows should be unity if the system achieves a steady state and the momentum source $\langle \rho S_z \rangle$ term averages to zero everywhere. Overall, we find that the $\langle P_{\text{tot}} \rangle / \langle W_{\text{tot}} \rangle$ ratio is typically within $\approx 20\%$ of unity for most heights $|z|$ up to $10h$, in agreement with hydrodynamic pressure gradients on

average roughly balancing the weight of the overlying gas. There is only a small dependence on the averaging scale ℓ_{patch} in the inner regions where there is only $\sim 1 - 10$ patches with $\ell_{\text{patch}} = 3$ kpc.

A notable systematic departure from $\langle P_{\text{tot}} \rangle / \langle W_{\text{tot}} \rangle \approx 1$ is found within a disk scale height in most panels: near the midplane, the hydrodynamic pressure is systematically below the weight of the overlying gas, as seen in Figure 7. Examining the approximations made in the balance equation derivation (§4), we identify a few possible contributions to these midplane “pressure dips.”

First, $\langle P_{\text{tot}} \rangle / \langle W_{\text{tot}} \rangle$ will only tend to unity when the momentum source term $\langle \rho S_z \rangle \approx 0$. In the limit of infinite resolution, the momentum from SNe and stellar winds in our subgrid model would be injected from point sources with zero net linear momentum by construction (Hopkins et al. 2018a). In practice, the momentum is injected over a set of ~ 32 neighboring resolution elements. In the presence of a vertical gradient in the distribution of momentum sources, the finite injection volume induces a non-zero $\langle \rho S_z \rangle$ contribution. A related effect arises in the subgrid model for radiation pressure, which includes a long-range component (Hopkins et al. 2019). The effects of smooth, large-scale radiation supporting the disk against gravity are also not included in P_{tot} because much of the momentum transferred from the

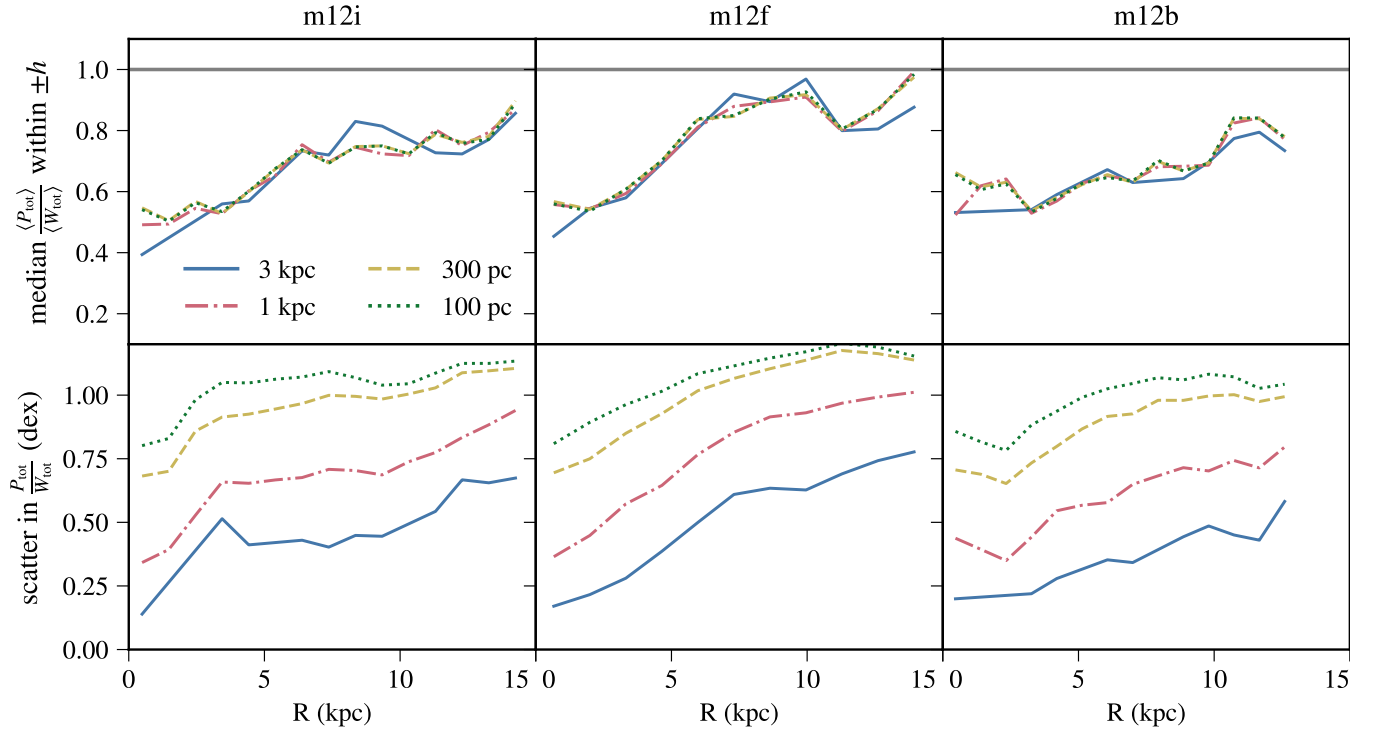


Figure 7. Radial profiles of the “height-median pressure ratio” (top) and the “height-median patch-to-patch scatter” (bottom) of the pressure ratio at redshift $z \sim 0$ and averaged over $\Delta t \approx 600$ Myr (approximately the ratio of the midplane pressure to the vertical weight of gas in the disk, see §4.2 for exact definitions). Both are evaluated using slabs within $\pm h$ of the midplane. Different patch sizes, ℓ_{patch} , are plotted in different line styles ranging from 3 kpc to 100 pc. Each column corresponds to a different simulation, indicated by the top headings. Overall, the height-median pressure systematically increases from ~ 0.5 to ~ 1 with increasing galactocentric radius, independent of the averaging scale. On the other hand, the patch-to-patch scatter increases both with radius and with decreasing averaging scale.

radiation field effectively reduces \bar{g} and is not captured by gas motions. Simple estimates indicate that the total momentum injected in the ISM by SNe is $\sim 5\times$ the momentum injected by direct radiation pressure (e.g., Faucher-Giguère et al. 2013). Thus, if we do not explicitly account for the momentum deposited in the disk by smooth radiation pressure, the ISM can appear “under-pressurized” by $\sim 20\%$ due to the missing radiation pressure. Second, the derivation in §4 neglected some correlations between the gas density and velocity fields, e.g. to set $\langle \rho u_R u_z \rangle = \langle u_R \rangle \langle \rho u_z \rangle = 0$ (Equation (9)), but these terms can be non-zero in general.

Finally, we note that we perform our analysis on cosmological zoom-in simulations, in which galaxies are still evolving during the analysis period and in which the galaxies can be perturbed by various time-dependent external forces (e.g., satellite galaxies). Thus, the galaxies are not in an exact statistical steady state and we expect larger departures from unity of the pressure ratio than in more idealized studies based on local boxes (or isolated galaxies), which can achieve equilibrium in a more controlled setting.

Overall, we find that on average vertical hydrodynamic pressure gradients in the simulations analyzed balance the weight of the overlying gas within tens of percent, out to $|z| \geq 10h$. In particular, the deviations from perfect average equilibrium are much smaller than the scatter in the KS relation ($\sim 0.5 - 1$ dex), which equilibrium models are sometimes invoked to explain (e.g., Faucher-Giguère et al. 2013; Hayward & Hopkins 2017; Orr et al. 2018, 2019b).

4.3 Pressure contributions from different phases and forms

Next, we analyze how the total ISM pressure is decomposed into a sum of partial pressures from different gas phases, as well as contributions from the thermal, velocity dispersion, and bulk flow pressure components defined in §3.2. We focus on radii $R \leq 9$ kpc, which are representative of the main parts of the galaxies (i.e. excluding the outermost, lower-density regions, where ISM pressure and star formation may be regulated differently, e.g. owing to the cosmic ultra-violet background) and a fiducial patch size $\ell_{\text{patch}} = 1$ kpc.

Figure 9 decomposes the total pressure into contributions from each gas phase and different components. From the top panel of Figure 9a, we see that the total pressure is provided nearly equally by partial pressures from the warm and cold phases within the galactic disk (i.e. within $\pm h$), mostly by the warm phase up to a height $|z| \leq 2.5h$, and by the hot phase for heights above that.

The bottom panel of Figure 9a shows the ratio between the local pressure in each gas phase $\bar{P}_{\text{tot}}^{\text{phase}}$ and the total ISM pressure P_{tot} . Since for each phase the partial and local pressures are related by a factor of the volume filling fraction, f_{phase} (see Equation (15)), the curves in the top and bottom panels for each phase are related by the z -dependent volume filling fraction.

We see that, for the most part, the phases are in roughly local total pressure equilibrium. As we discussed below, the different phases are in general however not in thermal pres-

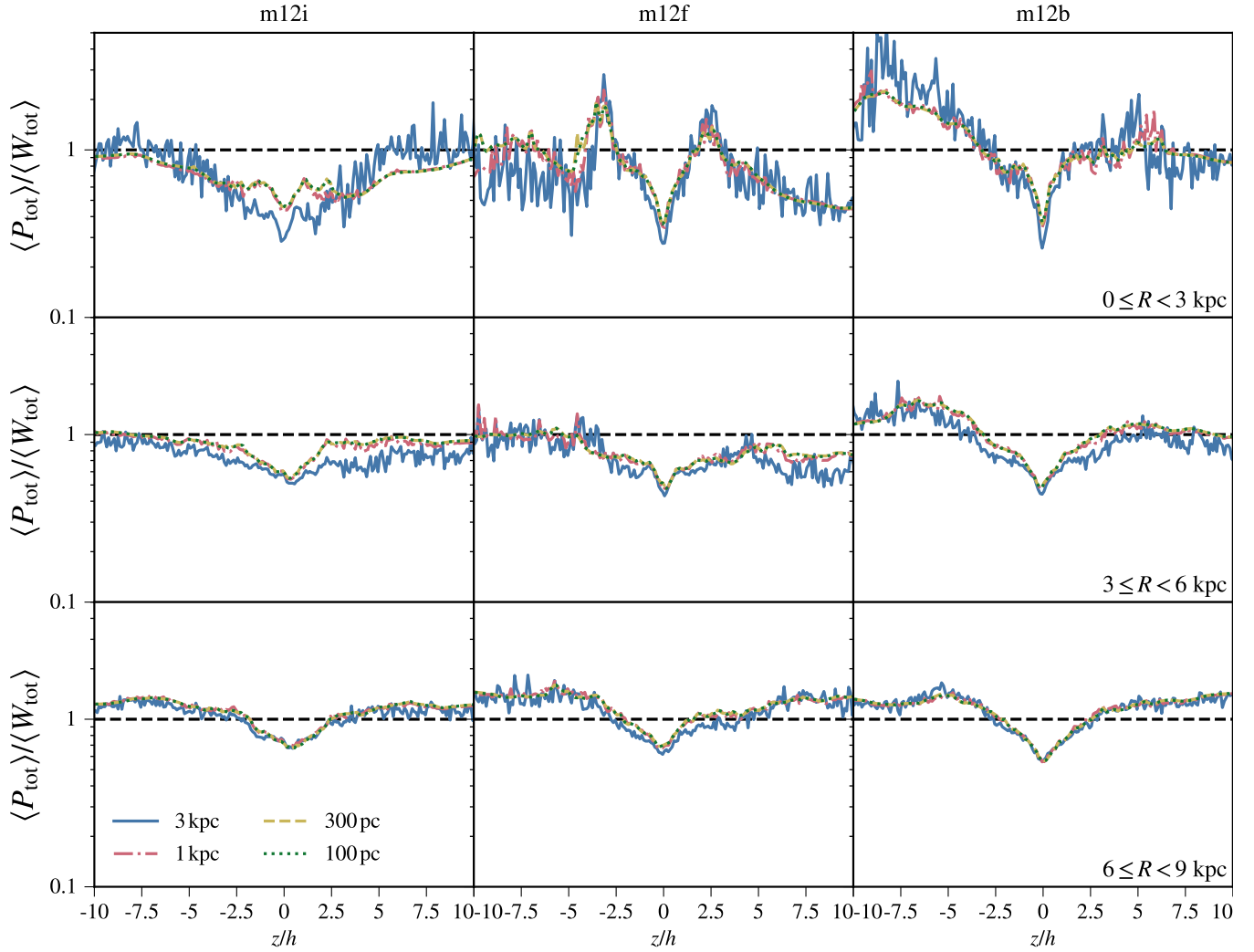


Figure 8. Vertical profiles of the annulus-averaged pressure-to-weight ratio at redshift $z \sim 0$ and averaged over the last $\Delta t \approx 600$ Myr as a function of vertical position z in units of the scale height h (see Figure 3 for scale height as a function of radius). Different patch sizes, ℓ_{patch} , are plotted in different line styles ranging from 3 kpc to 100 pc. Rows divide patches into radial annuli (denoted in the bottom right corner of the panels). In general, the pressure-to-weight ratio is ≈ 1 , representing approximate vertical pressure balance, with only small variation between different ℓ_{patch} . Systematic departures near the midplane are evident in all three simulations analyzed, indicating the importance of additional terms in equation (11) (see §4.2 for further discussion).

sure equilibrium owing to large non-thermal contributions. Near the midplane, the local pressure in the hot gas is typically larger than P_{tot} , which is expected for hot gas generated by SNe. This hot, shocked gas is initially overpressurized and drives outflows which escape the disk. The large scatter in the local pressures for all gas phases indicates that substantial local fluctuations in pressure balance between different phases are present. This is consistent with a dynamic ISM continuously reshaped by cooling, gravity, and feedback.

Figure 9b decomposes the total pressure in each gas phase into the different components P_{therm} , P_{disp} , and P_{bulk} as a function of height.⁹ Qualitatively, we find that the cold and warm phases have similar contributions from kinetic en-

ergy as a function of height, with P_{disp} dominating near the midplane and P_{bulk} becoming increasingly dominant with increasing height. As expected, the thermal pressure is negligible in the cold phase; it is also everywhere subdominant in the warm phase. The picture is markedly different for the hot phase, where P_{therm} dominates at all heights, followed by P_{bulk} , and with P_{disp} contributing only $\sim 2\%$ of the total pressure in the hot phase. The fact that the dispersion and bulk pressure components dominate in the cold and warm phases within the disk (i.e., within a few scale heights) indicates that rough equilibrium between *total* local pressures in different phases does not imply *thermal* pressure equilibrium between the phases. In particular, the thermal pressure

⁹ Note that, per the definition in Equation (15), the *fractional* partitioning between different forms of pressure is the same for

both the partial and local pressures, as the volume filling fraction cancels.

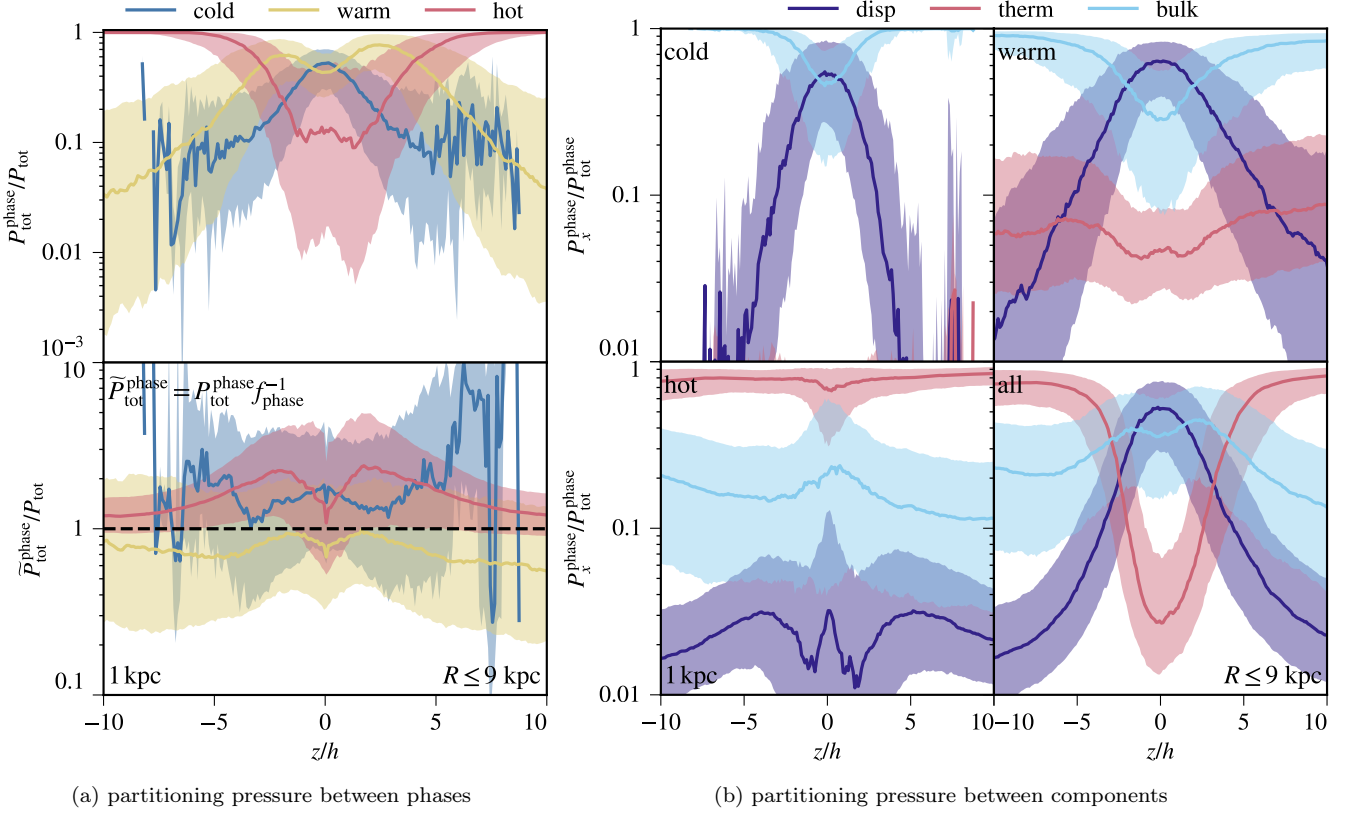


Figure 9. (a) *Top*: Partial pressures for the cold ($T \leq 10^3$ K), warm ($10^3 \leq T \leq 10^5$ K), and hot ($T \geq 10^5$ K) gas phases, normalized by the total ISM pressure P_{tot} as a function of height from the disk midplane. For each phase, all pressure components (thermal, dispersion, and bulk flow) are included. *Bottom*: “Local” pressure in each phase (\tilde{P} , defined in Equation (15), also including all components) relative to P_{tot} . These pressures differ from the partial pressures in the top panel by the inverse of the volume filling fraction f_{phase} . Ratios ~ 1 implies pressure equilibrium between phases. (b) Fractional contributions for each pressure component (thermal, dispersion, or bulk flow) to the total pressure in each gas phase (indicated at the bottom left of each panel). For the cold phase, the thermal contribution is so small that it is below the y -axis. In each panel, solid curves show patch-medians and shaded regions are the inter-quartile ranges, which quantify patch-to-patch scatter. These results are evaluated for m12i in patches with $R \leq 9$ kpc, using a fiducial patch size $\ell_{\text{patch}} = 1$ kpc, and in 10 snapshots spanning ~ 600 Myr at redshift $z \sim 0$. Considered together, (a) and (b) show that, the total ISM pressure is dominated by dispersion/turbulent pressure from the cold+warm gas near the midplane and by thermal pressure in the hot phase at large heights. This description of the ISM is independent of the averaging scale ℓ_{patch} , save for the partitioning of total “kinetic” pressure $P_{\text{kin}} \equiv P_{\text{disp}} + P_{\text{bulk}}$ between the dispersion/turbulence and bulk flow terms.

in the cold and warm phases are much lower than in the hot phase within the disk.¹⁰

The results shown in Figure 9 are independent of the averaging scale ℓ_{patch} , save for the partitioning of total “kinetic” pressure $P_{\text{kin}} \equiv P_{\text{disp}} + P_{\text{bulk}}$ between the dispersion/turbulence and bulk flow terms. Figure 10 shows how this partitioning varies with averaging scale and as a function of height from the midplane. As expected, the total kinetic pressure is independent of averaging scale, but the

relative importance of the dispersion and bulk flow contributions depends strongly on the averaging scale.

We expect a transition in the partitioning of P_{kin} at $\ell_{\text{patch}} \approx \ell_{\text{turb}}$, where ℓ_{turb} is the scale of the largest turbulent eddies. For $\ell_{\text{patch}} \gtrsim \ell_{\text{turb}}$, a single patch contains all scales relevant to the turbulent cascade. On the other hand, for $\ell_{\text{patch}} \lesssim \ell_{\text{turb}}$ the largest turbulent eddies cannot be contained in individual patches. In this limit, the larger turbulent eddies will appear as coherent bulk flows on the scale ℓ_{patch} , and will thus contribute to P_{bulk} instead of P_{disp} . Near the disk midplane, Figure 10 shows that P_{kin} transitions from being dominated by P_{disp} to being dominated by P_{bulk} as ℓ_{patch} decreases from 1 kpc to 300 pc. This indicates that the largest turbulent eddies (which contain most of the kinetic energy in a 3D turbulent cascade) have a scale $300 \text{ pc} \lesssim \ell_{\text{turb}} \lesssim 1 \text{ kpc}$. This is consistent with the idea that the largest turbulent eddies have a scale comparable to the scale height of the gas disk (see Figure 3).

¹⁰ Note that results regarding whether different phases are in thermal pressure balance with each other can depend on how the phases are defined. For example, Vijayan et al. (2020) define their warm phase as $5050 \leq T \leq 2 \times 10^4$ K and an “intermediate” phase as $2 \times 10^4 \leq T \leq 10^5$ K. Vijayan et al. find similar local thermal pressures in the warm and intermediate phases, but this is not necessarily inconsistent with our results since both of these phases defined by Vijayan et al. would be contained within our warm phase with $10^3 \leq T \leq 10^5$ K.

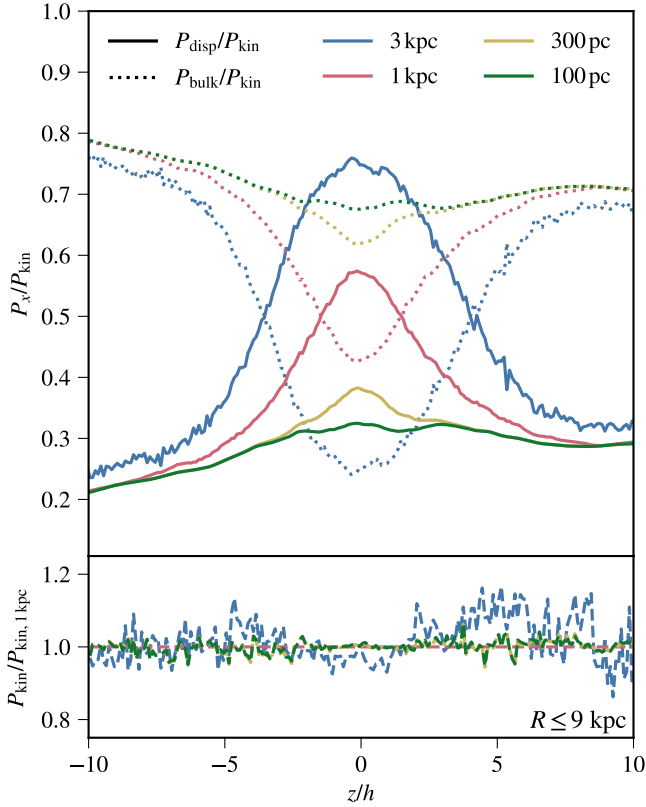


Figure 10. *Top:* analysis of how the partitioning of the total kinetic pressure $P_{\text{kin}} = P_{\text{disp}} + P_{\text{bulk}}$ (the sum of the dispersion/turbulent and bulk flow components) varies as a function of patch size (or averaging scale) ℓ_{patch} . The curves show patch-medians for m12i at redshift $z \sim 0$, over a period $\Delta t \approx 600$ Myr, and include all patches with $R \leq 9 \text{ kpc}$ using the fiducial $\ell_{\text{patch}} = 1 \text{ kpc}$. Curves of different color show what fractions of the total kinetic pressure are contributed by P_{disp} (solid) and P_{bulk} (dotted) for different values of ℓ_{patch} as a function of height. The partitioning of kinetic energy between P_{disp} and P_{bulk} varies strongly as a function of ℓ_{patch} and $|z|$. Near the midplane, the kinetic energy is concentrated in the velocity dispersion for large ℓ_{patch} and in bulk flows for small ℓ_{patch} . These limits correspond to whether ℓ_{patch} encompasses or is subsumed by individual turbulent eddies, suggesting that the largest turbulent eddies have a size scale between 1 kpc and 300 pc, comparable to the disk scale height. *Bottom:* The P_{kin} profiles relative to the values for $\ell_{\text{patch}} = 1 \text{ kpc}$ are plotted as dashed lines. These dashed curves show that the total kinetic pressure is constant with varying averaging scale (modulo stochastic fluctuations), which is expected by construction.

4.4 Vertical profiles of bulk velocities and velocity dispersions

Lastly, we analyze how the bulk velocity and the velocity dispersion of the gas in the different phases vary as a function of height from the midplane.

The top panel of Figure 11 shows how the bulk vertical velocities (corresponding to the slab-averaged u_z^s defined in §3.2) vary as a function of $|z|/h$, measured in 1 kpc patches. We see, for the cold gas, there is little to no vertical bulk motion within a few disk scale heights that is coherent throughout the disk. However, the significant fraction of

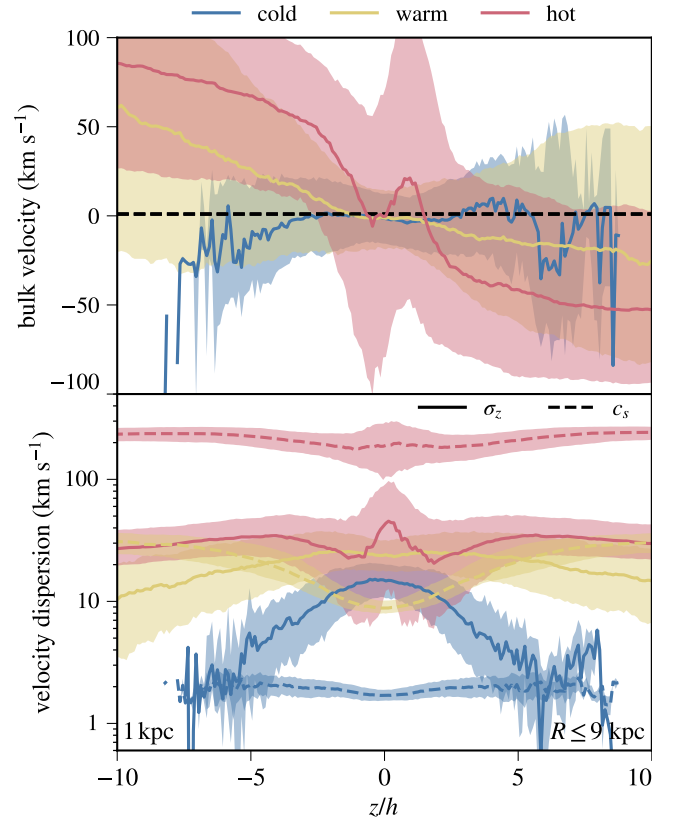


Figure 11. *Top:* vertical bulk velocity ($u_z^s = \langle u_z \rangle_M$) of each of the cold ($T \leq 10^3 \text{ K}$), warm ($10^3 \leq T \leq 10^5 \text{ K}$), and hot ($T \geq 10^5 \text{ K}$) gas phases. The solid curves show patch-medians for m12i in patches with $R \leq 9 \text{ kpc}$, using a fiducial patch size $\ell_{\text{patch}} = 1 \text{ kpc}$, and in 10 snapshots spanning ~ 600 Myr at redshift $z \sim 0$. Shaded regions are the interquartile range and are shown to quantify the patch-to-patch scatter. Positive values correspond to outflowing gas above the midplane ($z/h > 0$) and inflowing gas below the midplane (and vice versa for $\langle u_z \rangle_M < 0$). The hot gas is outflowing near the midplane and inflowing at larger heights. The same is true for the warm phase, on average, albeit at lower velocity. The cold phase, by contrast, has a median velocity ~ 0 at all heights with substantial, $\sim \pm 10 \text{ km s}^{-1}$, patch-to-patch scatter. *Bottom:* z -velocity dispersion and thermal sound speed for each gas phase. In both panels, quantities are evaluated for m12i in patches with $R \leq 9 \text{ kpc}$, using a fiducial patch size $\ell_{\text{patch}} = 1 \text{ kpc}$, and in 10 snapshots spanning ~ 600 Myr at redshift $z \sim 0$. By comparing the bulk velocities to the velocity dispersions, we see that the cold gas can be described as being dominated by supersonic turbulence in the disk, the warm gas by supersonic turbulence in the disk and transonic inflows/outflows at larger heights, and the hot gas by subsonic bulk inflows/outflows.

the total pressure contributed by P_{bulk} near the midplane in Figure 9b indicates that while the patch-median velocity is close to zero, the z -component of the cold gas velocity must fluctuate between positive and negative values between different analysis patches. There is some outflowing cold gas at $|z|/h \gtrsim 3$ but with large scatter in the bulk velocity due to the small amount of cold gas present at these heights. Similarly, there is little bulk motion in the warm gas near the midplane. Interestingly, both the warm gas and the hot gas are on average *inflowing* from $|z| \sim 10h$ to $|z| \sim 2h$, presumably corresponding to cooling of the hot atmospheres sur-

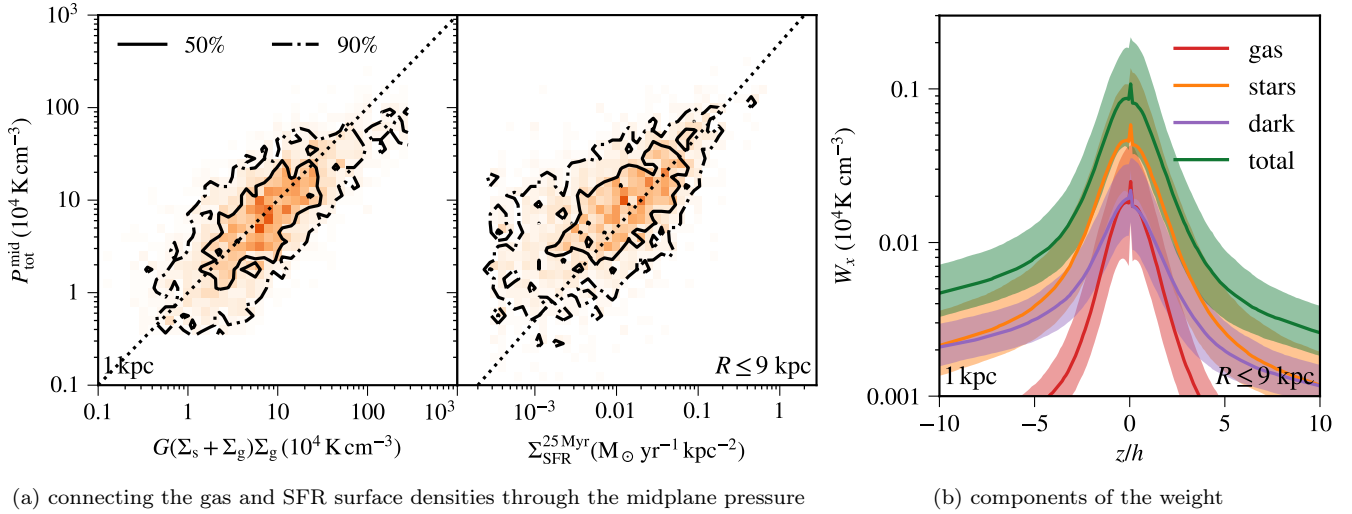


Figure 12. (a) Total midplane pressure (averaged within $|z| \leq h/8$) vs. (left) an analytic predictor for the weight of the disk gas ($G(\Sigma_s + \Sigma_g)\Sigma_g$) and (right) the star formation rate surface density averaged over 25 Myr. Contours are drawn containing 50% (solid) and 90% (dot-dashed) of the patches. A 1-1 line is plotted for reference in the left panel and a line corresponding to $P_{\text{tot}}^{\text{mid}} = \Sigma_{\text{SFR}} (1000 \text{ km s}^{-1})$ is plotted in the right panel. (b) Vertical profiles of the different gas weight components W_x , defined as in Equation (13) but in each case including only the contribution to the gravitational acceleration \bar{g} from the gas, stars, or dark matter. The total weight of the disk gas is also shown in green for comparison. Solid lines are the patch-medians and shaded regions show the inter-quartile range to give an indication of the patch-to-patch scatter. The weight from the stars dominates at all heights, consistent with the low gas mass fractions $f_g \sim 0.2 \ll 1$ (see Table 1). In all panels, quantities are evaluated for m12i in patches with $R \leq 9$ kpc, using a fiducial patch size $\ell_{\text{patch}} = 1$ kpc, and in 10 snapshots spanning ~ 600 Myr at redshift $z \sim 0$.

rounding the $\sim L^*$ galaxies we analyze (e.g., Stern et al. 2019; Hafen et al. 2020). Some of the inflowing gas can also correspond to the recycling of gas previously ejected by winds (Oppenheimer et al. 2010; Anglés-Alcázar et al. 2017).

Taking into account the full distribution of bulk velocities, all phases contribute to some degree to outflows. The median hot gas bulk velocity shows the most complex behavior. In and around the disk ($|z| \lesssim 2h$), the hot gas is primarily outflowing while at larger heights the bulk velocity changes sign, i.e. the hot gas becomes predominantly inflowing. Overall, the results are consistent with weak “fountain flows” in which winds leaving the disk do not reach far into the CGM before turning around or being incorporated into a circumgalactic cooling flow at larger heights. Such weak winds in low-redshift $\sim L^*$ galaxies contrast with the much faster and more highly mass-loaded outflows found around the high-redshift progenitors of the same galaxies in FIRE (Muratov et al. 2015; Stern et al., in prep.; Chan et al., in prep.).

The bottom panel of Figure 11 shows the z -velocity dispersion σ_z and thermal sound speed c_s as a function of height in each of the three different phases. By comparing the bulk velocities to the velocity dispersions in the top vs. bottom panels of the figure, we see that the cold gas can be described as being dominated by supersonic turbulence in the disk, the warm gas by supersonic turbulence in the disk and transonic inflows/outflows at larger heights, and the hot gas by subsonic bulk inflows/outflows.

The total velocity dispersion profiles ($\sim \max(c_s, \sigma_z)$) are roughly constant with height within the disk ($|z|/h \lesssim 2-3$).

5 DISCUSSION

5.1 The Toomre Q parameter

Self-regulation to marginal gravitational stability is a common assumption in models of galactic disks (e.g. Thompson et al. 2005; Krumholz & Burkert 2010; Faucher-Giguère et al. 2013; Hayward & Hopkins 2017; Orr et al. 2018). The “classic” single-component Q parameters for purely stellar or purely gaseous disks (Toomre 1964; Goldreich & Lynden-Bell 1965) are defined such that local axisymmetric perturbations in infinitely thin, homogeneous disks are gravitationally unstable when $Q \lesssim 1$. Marginal gravitational stability then predicts that disks should have $Q \approx 1$. For our adopted definition of Q_{eff} , which explicitly includes a stellar component and a gaseous component (see Equation (2)), the criterion for gravitational instability is slightly modified. In an extended stability analysis taking into account finite disk thickness and dissipation, Elmegreen (2011) showed that instability is expected in a two-component stars+gas disk when $Q_{\text{eff}} \lesssim 0.75$.

In §2 we noted that our disks have an effective Toomre Q_{eff} parameter that is nearly a constant ~ 3 throughout the main parts of the disk. At face value, since Q_{eff} is well above the nominal Elmegreen (2011) stability threshold, this suggests that the perturbations in the disks should be stable against gravitational collapse. This is however clearly not the case, since the disks are continuously forming stars and in the simulations, star formation is restricted to locally self-gravitating gas. This implies that even more detailed, multi-component analytic disk stability analyses such as that performed by Elmegreen (2011) are not in general sufficient to accurately model gravitational instabilities in realistic disk galaxies. This is not too surprising, since in

realistic galaxies the multiphase ISM is continuously stirred and compressed by both gravitational and stellar feedback processes, and dissipation occurs through a complex combination of gas cooling and kinetic dissipation. This leads to highly non-linear perturbations whose development cannot be accurately modeled by a linear analysis of a smooth disk model.

We note that the relatively large $Q_{\text{eff}} \sim 3$ that we find in the simulations goes in the direction that might be expected given the complexities just described. Namely, even though we find $Q_{\text{eff}} \sim 3$ when evaluating disk properties averaged in annuli, local gas densities (e.g. in spiral arms) can be much higher than average and thus lead to local gravitational instability and star formation. At the same time, large portions of the disk are “over-supported” (by e.g. super-bubbles and outflows) and do not form stars. In a different analysis of FIRE-2 simulations, Orr et al. (2019a) measured $Q \sim 1$ in ~ 1 kpc patches, but their analysis was focused on regions with non-zero SFR, whereas we include all gas in the present analysis. Although the analysis of Orr et al. (2019a) used a different expression for Q , we verified that this is a subdominant effect. Interestingly, Leroy et al. (2008) analyzed a sample of nearby, observed spiral galaxies and inferred values $Q_{\text{eff}} \sim 2 - 3$ consistent with those found in our simulations.

5.2 Vertical pressure balance as an explanation for the KS relation

As discussed in the introduction (§1), a basic assumption of equilibrium models for the KS relation is that galactic disks are in a state of approximate vertical hydrostatic balance, in the sense that the total ISM pressure near the midplane balances the weight of the overlying disk gas. The results of the previous section show that this requirement is satisfied for the long-lived, low-redshift, $\sim L^*$ simulated galactic disks analyzed in this paper. In particular, the results indicate that near the midplane most of the ISM pressure is provided by dispersive motions in the cold and warm phases, which in observations would typically be identified as supersonic or transonic turbulence. The equilibrium is quasi *hydrostatic* near the midplane in the sense that large-scale bulk flows are subdominant.

Vertical hydrostatic balance is however not sufficient on its own to explain the KS relation; for this, there must be a mechanism that couples Σ_{SFR} to Σ_g (or the gas surface density in a specific component, such as molecular gas). In feedback-regulated models, it is postulated that midplane ISM pressure is predominantly provided by momentum injection by stellar feedback. In this case, under simple assumptions for turbulence dissipation, the turbulent pressure $P_{\text{disp}} \sim \Sigma_{\text{SFR}}(p_*/m_*)$, where $p_*/m_* \sim 1000 \text{ km s}^{-1}$ is the momentum injected into the ISM by SNe¹¹ (e.g., Ostriker & Shetty 2011; Faucher-Giguère et al. 2013; Martizzi et al. 2015; Hayward & Hopkins 2017).

Figure 12a examines in more detail how, in the simulations analyzed, Σ_{SFR} and Σ_g are consistent with being coupled through the midplane pressure as in feedback-regulated, equilibrium KS models. In the left panel, we show the total midplane pressure $P_{\text{tot}}^{\text{mid}}$ in m12i versus the analytic estimate $G\Sigma_{\text{disk}}\Sigma_g = G(\Sigma_s + \Sigma_g)\Sigma_g$ for the weight of the overlying gas (such a weight-based expression is often how the midplane pressure is estimated in observations, e.g. Wong & Blitz 2002; Leroy et al. 2008; Herrera-Camus et al. 2017). In the right panel, we show how $P_{\text{tot}}^{\text{mid}}$ scales with Σ_{SFR} averaged over a timescale of 25 Myr (this roughly corresponds to the timescale needed for most of the feedback energy from SNe to be deposited into the ISM).

Consistent with the key results of §4 showing that the total midplane pressure is determined by the weight of the overlying gas, the left panel of Figure 12a shows a clear linear correlation, with approximately the same normalization, for $P_{\text{tot}}^{\text{mid}}$ vs. $G\Sigma_{\text{disk}}\Sigma_g$. On the right, we see that the midplane pressure also correlates with Σ_{SFR} , with a normalization $P_{\text{tot}}^{\text{mid}} \sim \Sigma_{\text{SFR}}(p_*/m_*)$ consistent with $p_*/m_* \sim 1,000 \text{ km s}^{-1}$. These scalings are qualitatively similar to those recently observed in Sun et al. (2020) (compare to their Figure 8). They also agree with the basic assumptions of feedback-regulated models and, when combined by eliminating the midplane pressure, imply a KS-type relation between Σ_{SFR} and Σ_g . KS-type relations are indeed realized in the FIRE simulations (Hopkins et al. 2014; Orr et al. 2018), though as discussed at length by Orr et al. (2018), a quantitative comparison with observations requires a detailed treatment of gas tracers.

The different contributions to the total gravitational acceleration in the disk are plotted in Figure 12b. The figure shows that, for the present simulations, the total gravitational acceleration is dominated by stars, implying that $P_{\text{tot}}^{\text{mid}} \propto G\Sigma_{\text{disk}}\Sigma_g \approx G\Sigma_*\Sigma_g$. In more gas-rich galaxies with gas mass fraction $f_g \sim 1$, however, the scaling can approach $P_{\text{tot}}^{\text{mid}} \propto G\Sigma_g^2$, which would result in a steeper scaling of Σ_{SFR} vs. Σ_g .

5.3 Relating the scatter in the pressure ratio to the observed scatter in the KS Relation

Observational work has shown that the scatter in the KS relation increases with decreasing averaging scale (Kennicutt et al. 2007; Bigiel et al. 2008, 2010; Leroy et al. 2008, 2017). It is tempting to suggest that deviations from pressure balance in our analysis, which also increase in magnitude as ℓ_{patch} is decreased, might drive (part of) the deviations from the mean KS relation.

In the simple feedback-regulated equilibrium framework, the equations from the previous section imply that

$$\Sigma_{\text{SFR}} \sim \left(\frac{P}{W}\right) \left(\frac{G}{p_*/m_*}\right) \Sigma_{\text{disk}}\Sigma_g, \quad (17)$$

where we omit the subscript ‘tot’ in the pressure-to-weight ratio P/W for simplicity.

Following a standard propagation-of-errors procedure for the scatter,

$$\left(\frac{\sigma_{\Sigma_{\text{SFR}}}}{\Sigma_{\text{SFR}}}\right)^2 = \left(\frac{\sigma_{P/W}}{P/W}\right)^2 + \left(\frac{\sigma_{\Sigma_{\text{disk}}}}{\Sigma_{\text{disk}}}\right)^2 + \left(\frac{\sigma_{\Sigma_g}}{\Sigma_g}\right)^2 + \text{o.t.}, \quad (18)$$

¹¹ There is a factor ~ 3 uncertainty in the correct normalization of p_*/m_* owing not only to uncertainties in the properties of SNe but also where they explode in the ISM and the degree to which the linear momentum from different SNe cancel in the disk.

where ‘o.t.’ are other terms that may contribute, e.g. finite sampling of star-forming regions (e.g., [Kruijssen & Longmore 2014](#)), time delays between star formation events and ISM pressure equilibration (e.g., [Orr et al. 2019b](#)), or variations in p_*/m_* . This implies that the observed fractional scatter in Σ_{SFR} (at fixed Σ_g) must be larger than the predicted fractional scatter in P/W for the model and simulations to be consistent with observations.

A crude comparison with available data on the scale-dependent scatter of the KS relation in local spiral galaxies similar to the simulated galaxies (e.g. in [Kennicutt et al. 2007](#); [Bigiel et al. 2010](#); [Leroy et al. 2017](#)) suggests that the scatter in P/W predicted by the simulations (see Figure 7) is consistent with, but not negligible compared to, the total fractional scatter in observationally-inferred Σ_{SFR} . This indicates that time- and space-dependent scatter in the pressure-to-weight ratio can potentially account for a substantial portion of the scatter in the KS relation. Thus, it would be interesting for future work to model in more detail how fluctuations in P/W translate into scatter in the observed KS relation.

5.4 Previous studies of vertical pressure balance

Our results paint a dynamic picture of the ISM, one where individual patches can instantaneously be out of vertical pressure balance, but after averaging over significant portions of the disk there is an average pressure-to-weight ratio of ≈ 1 . The time-averaged equilibrium between the total ISM pressure and the weight of the overlying gas is broadly consistent with what has been found in previous studies of vertical hydrostatic balance in more idealized simulations of Milky Way-like conditions based on local boxes or non-cosmological simulations of isolated galaxies (e.g., [Kim et al. 2013](#); [Vijayan et al. 2020](#); [Benincasa et al. 2016](#); [Benincasa et al. 2019b](#)). Our finding that the total ISM pressure near the midplane is predominantly provided by dispersion/turbulence in the cold and warm phases is also consistent with previous simulations (although the cold phase was not always explicitly modeled), as is the fact that the hot gas generated by SNe tends to be over-pressurized in the disk and drive outflows (e.g., [Vijayan et al. 2020](#)).

One interesting difference between our results and the local-box simulation of [Vijayan et al. \(2020\)](#), which goes in a counter-intuitive direction, concerns the variability of the SFR in 1 kpc^2 patches (the size of the box in [Vijayan et al.](#)). Despite the idealized boundary conditions, the SFR in [Vijayan et al.](#)’s solar neighborhood local box experiences bursts with peak-to-trough amplitude exceeding an order of magnitude. In our cosmological simulations, the SFR similar in 1 kpc^2 patches varies by only a factor ~ 2 even though there are in principle many additional sources of perturbation (e.g. gas flows from different parts of the disk, spiral structure, or inflows/outflows/substructure in the galactic halo). One possible explanation for this difference in SFR burstiness is that the main form of stellar feedback in [Vijayan et al.](#)’s TIGRESS simulations is SNe ([Kim & Ostriker 2018](#)), whereas the FIRE-2 simulations also include photoionization and radiation pressure. [Hopkins et al. \(2019\)](#) showed that including these radiative feedback processes smooths galaxy-scale SFRs in time. In particular, ‘early’ radiative feedback limits the star formation efficiency of individual molecular clouds,

which in turn limits the formation of large clusters of SNe exploding at the same time.

We can also compare to recent observational studies that have sought to address whether galaxies achieve a state of dynamical equilibrium. Recently, [Sun et al. \(2020\)](#) carried out a detailed observational analysis of this question for 28 nearby star-forming galaxies and focusing on the cold molecular gas. Overall, they find very good agreement between the total pressure in the cold gas (dominated by dispersive/turbulent motions, as in our simulations) and the total weight of the overlying gas. [Sun et al. \(2020\)](#) note that in the observations, self-gravity of the gas is important to verify dynamical equilibrium in the densest cold gas. In our simulation analysis, self-gravity is implicitly neglected for substructures on scales $\ll \ell_{\text{patch}}$, but this is not a large effect because we exclude gas above a density cut of 50 cm^{-3} from our results (though not as a source of gravity, for which all mass is considered; see §3.3). Moreover, as summarized in §2.1, self-gravitating gas rapidly turns into stars in the simulations.

6 SUMMARY & CONCLUSIONS

We use the FIRE-2 galaxy formation simulations to investigate the structure and properties of the ISM in the disks of $z \sim 0$ Milky Way-mass galaxies. We analyze in particular the degree to which the ISM pressure can be modeled as being in quasi-hydrostatic balance with the weight of the overlying gas, an assumption which is the basis of many models for the structure of disk galaxies (e.g., [Boulares & Cox 1990](#)) as well as models of the origin of the Kennicutt-Schmidt relation (e.g., [Thompson et al. 2005](#); [Ostriker & Shetty 2011](#); [Faucher-Giguère et al. 2013](#); [Hayward & Hopkins 2017](#); [Orr et al. 2018](#)). We also analyze how the total ISM pressure is partitioned between different gas phases and different forms of pressure (thermal vs. dispersion/turbulence vs. bulk flows). This allows us to test the assumption of pressure balance between different phases, which is central to classic models of the ISM (e.g., [McKee & Ostriker 1977](#)).

Relative to previous analyses of ISM pressure balance in more idealized local-box simulations or in simulations of isolated galaxies (e.g., [Kim et al. 2013](#); [Benincasa et al. 2016](#); [Vijayan et al. 2020](#)), our study improves on previous work by including the full cosmological context. This includes inflows and outflows in the circumgalactic medium, as well as potential dynamical disturbances from satellite galaxies. Moreover, the FIRE-2 simulations implement a more complete set of stellar feedback processes, including Type II and Ia SNe, stellar winds, photoelectric heating, photoionization, and radiation pressure.

On average, we find that the total ISM pressure within the galactic disks approximately balances the weight of the overlying gas, as expected for systems in quasi-hydrostatic equilibrium. In more detail, the simulated galaxies exhibit deviations from exact vertical pressure balance with a scatter that increases with decreasing averaging scale (see Figure 7). By breaking the ISM pressure into different gas phases and forms, we find that the total ISM pressure in the disk is dominated by kinetic pressure due to velocity dispersion in the warm and cold phases.

Observationally, these dominant forms of pressure may

be identified as transonic and supersonic turbulence in the warm ionized medium (WIM), the warm neutral medium (WNM), and the cold neutral medium (CNM). This turbulence likely has important implications for models of the thermal instability and chemistry in the ISM (e.g. [Saury et al. 2014](#); [Bialy et al. 2019](#)). Above a few disk scale heights, the gas pressure is dominated instead by thermal pressure in the hot phase which drives outflows. In our simulations of low-redshift, $\sim L^*$ galaxies these outflows are weak and may be described as “fountain flows.”

We find that while the different gas phases are in rough *total* pressure equilibrium with each other within the disk, they are not in *thermal* pressure balance. In particular, since velocity dispersion pressure dominates in the warm and cold phases, these would appear to be highly under-pressurized if only their thermal pressure were considered.

A novel aspect of our analysis is the explicit decomposition of the total kinetic pressure between “dispersion” and “bulk flow” components (see §3.2). While total kinetic pressure is independent of averaging scale, the partitioning between the dispersion and bulk flow components depends on the scale of random motions, which can correspond to turbulent eddies, compared to the patch size. In the disk, we find the bulk flow term becomes dominant over the dispersion term as the averaging scale is decreased from 1 kpc to 300 pc (see Figure 10). This is consistent with the scale of the largest turbulent eddies being comparable to the ISM scale height.

We also quantify the contributions from different galactic components to the total gravitational acceleration. For the Milky Way-mass galaxies analyzed here, which have modest gas fractions $f_g \sim 0.2$, the gravitational acceleration due to stars is the most important near the midplane, while the gravity due to the dark matter and the self-gravity of the gas are strongly subdominant. Correspondingly, we find that the simple analytic expression $P_{\text{tot}}^{\text{mid}} \sim G(\Sigma_s + \Sigma_g)\Sigma_g$ predicts the total midplane pressure reasonably well.

Our analysis also shows that the ISM pressure scales linearly with the star formation rate surface density, $P_{\text{tot}}^{\text{mid}} \propto \Sigma_{\text{SFR}}$. Taken in concert, these results can explain why galaxies obey KS-type relations of the form $\Sigma_{\text{SFR}} \propto \Sigma_g^b$ (for exponents $b \approx 1 - 2$ which depend on how different mass components contribute to gravity, as well as the gas tracer used) as a result of vertical pressure balance in the ISM of disk galaxies. If this is correct, the scale-dependent fluctuations we find in vertical pressure balance provide insight into one source of scatter in the observed KS relation.

In the future, it will be important to explore the dynamical balance not only in stable Milky Way-like disk galaxies but also in the ISM of galaxies that exhibit highly bursty star formation, such as at high redshift. It will also be important to investigate whether the simulations predict a well defined KS relation in such regimes, and if so, how it might be realized if vertical pressure balance is not sustained. Additionally, the simulations analyzed in this paper neglected some potentially important physics, such as magnetic fields and cosmic rays. The effects of these additional processes on ISM balance will be explored in a separate study (Chan et al., in prep.).

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REFERENCES

- Anglés-Alcázar D., Faucher-Giguère C.-A., Kereš D., Hopkins P. F., Quataert E., Murray N., 2017, *MNRAS*, **470**, 4698
- Benincasa S. M., Wadsley J., Couchman H. M. P., Keller B. W., 2016, *MNRAS*, **462**, 3053
- Benincasa S. M., et al., 2019a, arXiv e-prints, [p. arXiv:1911.05251](#)
- Benincasa S. M., Wadsley J. W., Couchman H. M. P., Pettitt A. R., Keller B. W., Woods R. M., Grond J. J., 2019b, arXiv e-prints, [p. arXiv:1911.05252](#)
- Bialy S., Neufeld D., Wolfire M., Sternberg A., Burkhart B., 2019, *ApJ*, **885**, 109
- Bigiel F., Leroy A., Walter F., Brinks E., de Blok W. J. G., Madore B., Thornley M. D., 2008, *AJ*, **136**, 2846

- Bigiel F., Walter F., Blitz L., Brinks E., de Blok W. J. G., Madore B., 2010, *AJ*, 140, 1194
- Bland-Hawthorn J., Gerhard O., 2016, *ARA&A*, 54, 529
- Boulares A., Cox D. P., 1990, *ApJ*, 365, 544
- Bryan G. L., Norman M. L., 1998, *ApJ*, 495, 80
- Elmegreen B. G., 2011, *ApJ*, 737, 10
- Elmegreen B. G., Scalo J., 2004, *ARA&A*, 42, 211
- Faucher-Giguère C.-A., 2018, *MNRAS*, 473, 3717
- Faucher-Giguère C.-A., Quataert E., Hopkins P. F., 2013, *MNRAS*, 433, 1970
- Field G. B., Goldsmith D. W., Habing H. J., 1969, *ApJ*, 155, L149
- Fielding D., Quataert E., Martizzi D., Faucher-Giguère C.-A., 2017, *MNRAS*, 470, L39
- Fielding D., Quataert E., Martizzi D., 2018, *MNRAS*, 481, 3325
- Gallagher M. J., et al., 2018, *ApJ*, 858, 90
- Garrison-Kimmel S., et al., 2017, *MNRAS*, 471, 1709
- Garrison-Kimmel S., et al., 2018, *MNRAS*, 481, 4133
- Garrison-Kimmel S., et al., 2019, *MNRAS*, 487, 1380
- Genzel R., et al., 2010, *MNRAS*, 407, 2091
- Gill S. P. D., Knebe A., Gibson B. K., 2004, *MNRAS*, 351, 399
- Goldreich P., Lynden-Bell D., 1965, *MNRAS*, 130, 125
- Guszejnov D., Grudić M. Y., Offner S. S. R., Boylan-Kolchin M., Faucher-Giguère C.-A., Wetzel A., Benincasa S. M., Loebman S., 2020, *MNRAS*, 492, 488
- Hafen Z., et al., 2020, *MNRAS*, 494, 3581
- Hani M. H., Hayward C. C., Orr M. E., Ellison S. L., Torrey P., Murray N., Wetzel A., Faucher-Giguère C.-A., 2020, *MNRAS*, 493, L87
- Hayward C. C., Hopkins P. F., 2017, *MNRAS*, 465, 1682
- Heckman T. M., Thompson T. A., 2017, Galactic Winds and the Role Played by Massive Stars. p. 2431, doi:10.1007/978-3-319-21846-5_23
- Hennebelle P., Falgarone E., 2012, *A&ARv*, 20, 55
- Herrera-Camus R., et al., 2017, *ApJ*, 835, 201
- Hopkins P. F., 2015, *MNRAS*, 450, 53
- Hopkins P. F., Quataert E., Murray N., 2011, *MNRAS*, 417, 950
- Hopkins P. F., Narayanan D., Murray N., 2013, *MNRAS*, 432, 2647
- Hopkins P. F., Kereš D., Oñorbe J., Faucher-Giguère C.-A., Quataert E., Murray N., Bullock J. S., 2014, *MNRAS*, 445, 581
- Hopkins P. F., et al., 2018a, *MNRAS*, 477, 1578
- Hopkins P. F., et al., 2018b, *MNRAS*, 480, 800
- Hopkins P. F., Grudić M. Y., Wetzel A., Kereš D., Faucher-Giguère C.-A., Ma X., Murray N., Butcher N., 2019, *MNRAS*
- Hopkins P. F., Grudić M. Y., Wetzel A., Kereš D., Faucher-Giguère C.-A., Ma X., Murray N., Butcher N., 2020, *MNRAS*, 491, 3702
- Kennicutt Jr. R. C., 1998, *ApJ*, 498, 541
- Kennicutt Jr. R. C., Evans N. J., 2012, *Annual Review of Astronomy and Astrophysics*, 50, 531
- Kennicutt Jr. R. C., et al., 2007, *ApJ*, 671, 333
- Kim C. G., Ostriker E. C., 2015, *ApJ*, 815, 67
- Kim C.-G., Ostriker E. C., 2018, *ApJ*, 853, 173
- Kim C.-G., Kim W.-T., Ostriker E. C., 2011, *ApJ*, 743, 25
- Kim C.-G., Ostriker E. C., Kim W.-T., 2013, *ApJ*, 776, 1
- Knollmann S. R., Knebe A., 2009, *ApJ*, 182, 608
- Kruijssen J. M. D., Longmore S. N., 2014, *MNRAS*, 439, 3239
- Krumholz M. R., 2014, *Phys. Rep.*, 539, 49
- Krumholz M., Burkert A., 2010, *ApJ*, 724, 895
- Krumholz M. R., Burkert B., Forbes J. C., Crocker R. M., 2018, *MNRAS*, 477, 2716
- Leitherer C., et al., 1999, *ApJ*, 123, 3
- Leroy A. K., Walter F., Brinks E., Bigiel F., de Blok W. J. G., Madore B., Thornley M. D., 2008, *AJ*, 136, 2782
- Leroy A. K., et al., 2017, *ApJ*, 846, 71
- Licquia T. C., Newman J. A., Brinchmann J., 2015, *ApJ*, 809, 96
- Ma X., Hopkins P. F., Faucher-Giguère C.-A., Zolman N., Muratov A. L., Kereš D., Quataert E., 2016, *MNRAS*, 456, 2140
- Mac Low M.-M., Klessen R. S., 2004, *Reviews of Modern Physics*, 76, 125
- Martizzi D., Faucher-Giguère C.-A., Quataert E., 2015, *MNRAS*, 450, 504
- Martizzi D., Fielding D., Faucher-Giguère C.-A., Quataert E., 2016, *MNRAS*, 459, 2311
- McKee C. F., Ostriker J. P., 1977, *ApJ*, 218, 148
- McKee C. F., Ostriker E. C., 2007, *ARA&A*, 45, 565
- Muratov A. L., Kereš D., Faucher-Giguère C.-A., Hopkins P. F., Quataert E., Murray N., 2015, *MNRAS*, 454, 2691
- Oppenheimer B. D., Davé R., Kereš D., Fardal M., Katz N., Kollmeier J. A., Weinberg D. H., 2010, *MNRAS*, 406, 2325
- Orr M. E., et al., 2018, *MNRAS*, 478, 3653
- Orr M. E., et al., 2019a, arXiv e-prints, p. arXiv:1911.00020
- Orr M. E., Hayward C. C., Hopkins P. F., 2019b, *MNRAS*, 486, 4724
- Ostriker E. C., Shetty R., 2011, *ApJ*, 731, 41
- Ostriker E. C., McKee C. F., Leroy A. K., 2010, *ApJ*, 721, 975
- Saintonge A., et al., 2011, *MNRAS*, 415, 32
- Saury E., Miville-Deschênes M. A., Hennebelle P., Audit E., Schmidt W., 2014, *A&A*, 567, A16
- Shetty R., Ostriker E. C., 2012, *ApJ*, 754, 2
- Sparre M., Hayward C. C., Feldmann R., Faucher-Giguère C.-A., Muratov A. L., Kereš D., Hopkins P. F., 2017, *MNRAS*, 466, 88
- Spitzer Lyman J., 1956, *ApJ*, 124, 20
- Steidel C. C., Erb D. K., Shapley A. E., Pettini M., Reddy N., Bogosavljević M., Rudie G. C., Rakic O., 2010, *ApJ*, 717, 289
- Stern J., Fielding D., Faucher-Giguère C.-A., Quataert E., 2019, *MNRAS*, 488, 2549
- Sun J., et al., 2020, *ApJ*, 892, 148
- Tacconi L. J., Genzel R., Sternberg A., 2020, arXiv e-prints, p. arXiv:2003.06245
- Thompson T. A., Quataert E., Murray N., 2005, *ApJ*, 630, 167
- Toomre A., 1964, *ApJ*, 139, 1217
- Torrey P., Hopkins P. F., Faucher-Giguère C.-A., Vogelsberger M., Quataert E., Kereš D., Murray N., 2017, *MNRAS*, 467, 2301
- Vijayan A., Kim C.-G., Armillotta L., Ostriker E. C., Li M., 2020, *ApJ*, 894, 12
- Wang B., Silk J., 1994, *ApJ*, 427, 759
- Wellons S., Faucher-Giguère C.-A., Anglés-Alcázar D., Hayward C. C., Feldmann R., Hopkins P. F., Kereš D., 2020, arXiv e-prints, p. arXiv:1908.05274
- Wetzel A. R., Hopkins P. F., Kim J.-h., Faucher-Giguère C.-A., Kereš D., Quataert E., 2016, *ApJ*, 827, L23
- Wong T., Blitz L., 2002, *ApJ*, 569, 157

APPENDIX A: DEPENDENCE OF RESULTS ON ANALYSIS PARAMETERS

In Figure A1 we present two tests of how our results depend on the time sampling of snapshots in the simulations. The reference result corresponds to the vertical pressure-to-weight profile for the m12i simulation presented in the main text for $\ell_{\text{patch}} = 1$ kpc and $3 \leq R \leq 6$ kpc (see Figure 8). For the first test, we split the 10 snapshots used in the main analysis in two halves, each spanning 300 Myr (the m12i-first half and m12i-second half curves in the figure). For the second test, we use $3\times$ as many snapshots over the same total 600 Myr analysis period, with a spacing of ≈ 20 Myr between snapshots (m12i-triple curve). We find that our main results regarding the pressure-to-weight profile are insensi-

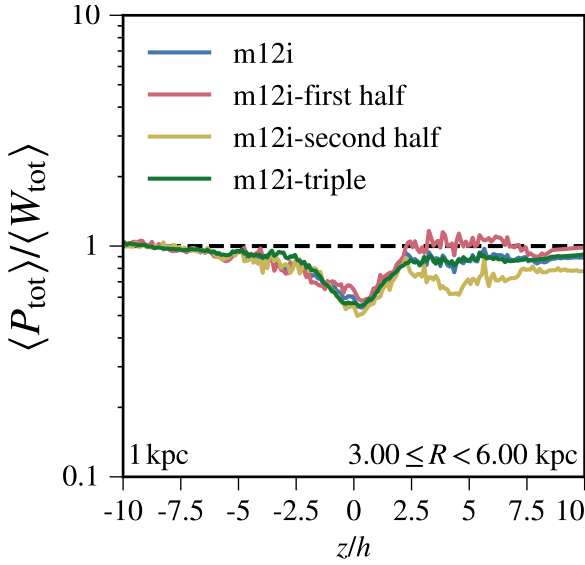


Figure A1. Vertical profiles of the annulus-averaged pressure ratio (directly analogous to Figure 8) for two different methods of time averaging in patches with $3 \leq R \leq 6$ kpc, using a fiducial patch size $\ell_{\text{patch}} = 1$ kpc. In the first case, the 10 fiducial snapshots used in the main analysis are split in half, each spanning 300 Myr of the fiducial 600 Myr period (m12i-first half and m12i-second half curves). In the second, $3\times$ as many snapshots are used to span the fiducial 600 Myr period with a spacing of ≈ 20 Myr between snapshots. A reference curve corresponding to the main analysis is provided for comparison (m12i).

tive to the number of snapshots in the time average and the period of time over which the average was performed. The results of the same tests for different radial bins are similar.

In Figure A2 we test the dependence of the results on the analysis slab thickness ($\delta z = h/50, h/10$, and $h/2$). We conclude again that the pressure-to-weight profiles are well converged for the fiducial slab thickness $\delta z = h/10$ at all radii and heights considered.

APPENDIX B: DEPOSITION OF MFM RESOLUTION ELEMENTS ONTO ANALYSIS GRID

Here we describe the method by which we smoothly deposit portions of MFM resolution elements that overlap multiple analysis slabs into those slabs (see Figure 4). The procedure is analogous to standard gridding of smooth particle hydrodynamics (SPH) particle data.

First, we determine whether a given MFM resolution element indexed by i overlaps multiple slabs by comparing its smoothing length h_{sml}^i to the grid spacing both horizontally (ℓ_{patch}) and vertically (δz). For this purpose, we consider an MFM resolution element to overlap with multiple slabs if and only if

$$h_{\text{sml}}^i > l_{\text{ovlp}} \equiv \max\left(\frac{\sqrt{3}}{2}\ell_{\text{patch}}, \delta z\right). \quad (\text{B1})$$

If it does, we assign fractions of the MFM resolution element's mass to overlapping slabs using a cubic spline kernel w such that the mass that resolution element i contributes

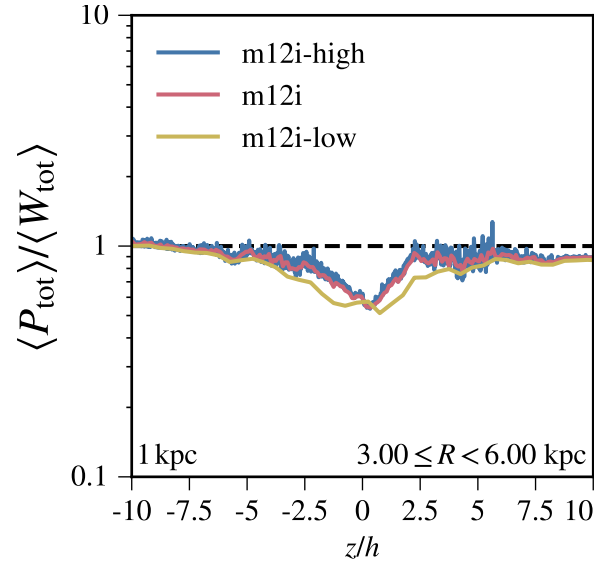


Figure A2. Similar to Figure A1 but for different analysis slab thicknesses: $0.2\times$, $1\times$, and $5\times$ the fiducial slab thickness of $h/10$ (m12i-high, m12i, and m12i-low respectively). The profiles results are well converged for the fiducial slab thickness $\delta z = h/10$ at all radii and heights.

to slab s is

$$m_{is} = \begin{cases} m^i \frac{w(r_{is}, h_{\text{sml}}^i) V_{\text{sml}}}{\sum_j w(r_{js}, h_{\text{sml}}^i) V^j} & h_{\text{sml}}^i > l_{\text{ovlp}} \\ m^i & h_{\text{sml}}^i \leq l_{\text{ovlp}} \end{cases}, \quad (\text{B2})$$

where j sums over the slabs whose centroid the MFM element's smoothing kernel intersects, r_{js} the distance from the resolution element centroid to the slab center, and $V^j = m_j/\rho_j$ is the effective volume of the MFM element. The sum in the denominator renormalizes the mass deposited to ensure mass conservation. Slabs are then assigned velocity components that are appropriate mass-weighted averages given contributions of overlapping MFM resolution elements.

The expression for the cubic spline kernel is:

$$w(r, h_{\text{sml}}) = \frac{8}{\pi h_{\text{sml}}^3} \begin{cases} 1 - 6(r/h_{\text{sml}})^2 + 6(r/h_{\text{sml}})^3 & 0 \leq r/h_{\text{sml}} < 0.5 \\ 2(1 - r/h_{\text{sml}})^3 & 0.5 \leq r/h_{\text{sml}} < 1 \\ 0 & 1 < r/h_{\text{sml}} \end{cases}. \quad (\text{B3})$$

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